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THE BIG QUESTION

ARE WE ALONE IN THE UNIVERSE? HOW COMMON IS LIFE?

NAI/NExSS Team: Mission to Young Earth 2.0

How To Find a Planet Habitable With Life?

• *How did life form and evolve on early Earth? Search for life in the Universe starts at our own planet and the Solar System.* • *What is the role of a host star on the origin and evolution of life on exoplanets?* • *How can life be detected ?*

• *30% fainter* $300/6$

- *Earth and Mars are icy balls* CO2 mixing ratio of 300 ppmv and a fixed relative humidity were assumed in the calculation. (From Kasting, Toon & Pollack 1988). completely frozen surfaces of Earth and Marshall and problem with habitability of young Earth and Marshall the young Earth and Marshall the second with the second of young Earth and Marshall the second with the second of the second with the second with the second with the se
- *Earth and Mars should not be habitable*
- *Faint Young Sun paradox*

USNO Scientific Colloquium, 01/26/17

Te, computed by a one-dimensional, radiative-convective climate model. The upper dashed curve represents the calculated mean global surface temperature, *Ts*. A fixed

Longitudinal surface magnetic field from Stokes V and I LSD profiles for a number of *high-g sepctral lines*

 Magnetism of The Young Sun

Marsden et al. 2014; Vidotto et al. 2014

*Correlation between age t and rotation period Prot for the stars, indicating that the non*accreting stars follow the Skumanich law (t ~ P_{rot}^2) \overline{P}_{rot}

USNO Scientific Colloquium, 01/26/17 **Typically about 10% of CME** kinetic energy goes into SEPs **Similar to flare energy Expect Ground Level Enhancement events (GLEs) are associated with energetic CMEs**

Carrington Event, Sept 1-2, 1859

Extreme CMEs:

 $E_{flare} = 2 \times 10^{33} \text{ erg} = 0.1 \text{ E}_{cme}$

• *Giant spots, flare and CME event* • *Most severe SW event in history* • *Lasted Eight days long* • *Aurorae at equatorial latitudes* • *Global telegraph network disrupted, operators suffered electric shocks* • *Magnetometers driven off scale* • *Energy in CME* \sim 2 x 10³³ erg • **Frequency – 250 events/d 4 Gyr ago** • *~ 1 per 300 yr today!* • *What to expect?*

Launch 2007 Kepler

 Surprises from Kepler Mission

Photometry: stellar brightness changes caused by transiting terrestrial planets

Variability range in mmag

Monitoring 150,000 stars at 30 min cadence for 4 years!

Over 4000 planet candidates

And a very surprising discovery of superflares on host stars

Paleo Space Weather: Dense & Fast Solar Wind and XUV flux

1 Rsun to 3 AU (Airapetian & Usmanov 2016)

X5 flare dominated XEUV from young Suns and red dwarfs Airapetian et al. 2017

Sources vs Sinks

Habitability of Early Earth:

Atmospheric Pressure

Outgassing (volcanic+tectonic) Solar Weather

Michail & Sverjensky Nature 2014 Airapetian et al. Nature Geoscience 2016

Solar wind, XUV, Precipitating particles SEP

Early Earth: Solar Energetic Particle Events

Prospects For Life on Early Earth & Titan

The dominant molecule Extremely hard to dissociate Triple bonds: 10 eV/atom vs 5.2 *eV*/*atom for* O_2

Prebiotic chemistry needs to break N_2 -> 2 N : • *VUV emission at λ < 100 nm (early Sun, M dwarfts)* • *Lightning discharge* • *Energetic particles (e, p)*

• $CO₂$ alone does not solve the problem \boldsymbol{p} robern 1.8 Ga, so the so that \boldsymbol{p} 100 increase), while fCH4 and fCO2 were held constant at their present values. $\mathcal{C}(\mathcal{O})$ with $\mathcal{O}(\mathcal{O})$ ing to our model, an increase in N2O by a factor of 100 would og vot solve the 1 to hou souve lite t O2 concentration did not reach the present atmospheric level

Other parameters are the same as in Fig. 2.

• $N_2\overline{O}$ is 300 x more potent that CO_2 interest here. Nevertheless, extremely low CO2 levels in the \mathcal{A} and \mathcal{A} Mid-Proterozoic, as well. The Rosing et al. argument has $\mathbf{N} \cap \mathbf{S}$ of \mathbf{S} \mathbf{N} assumed for \mathbf{N} as \mathbf{N} of \mathbf{N} of \mathbf{N} of \mathbf{N} lar to that from \mathbb{Z}/N , at identical concentrations, \mathbb{Z}/N \mathbf{u} and \mathbf{v} and \mathbf{v} (see, e.g. Berkonic \mathbf{v}) Marshall, 1964; Knoll, 1979; Canfield & Teske, 1996). The atmospheric lifetime of methane does not depend strongly on

we varied n2O concentrations from 0.3 to 30 ppm (a factor of a factor of a factor of a factor of a factor of a
The 30 ppm (a factor of a factor of a

Because the absorption bands of CH4 and N2O overlap, we

One of the uncertainties in this climate calculation is

calculated their combined effect explicitly (Fig. 3B).

- *Shielded by SO₂ and H₂S (240 nm)* \cdots is limited that side \cdots ited by the supply of reductant, i.e. organic matter, and not by the thermodynamics of carbonate formation (Dauphas & a stronger greenhouse gas when added to a putative Proterozoic atmosphere, but in terms of climate sensitivity (expressed in degrees was done that if \mathcal{O}_p $\sum_{i=1}^{\infty}$ in $\sum_{i=1}^{\infty}$ (a.10 and H.N.I Paul $\frac{1}{2}$ where $\frac{1}{2}$ $\frac{1}{2}$ $\frac{1}{2}$ would have dependent
- *FYS paradox can be resolved at* > 10 ppmv! $\delta t > U$ limit is invalid. \mathbf{w} discount the parameters on \mathbf{w} N2O are equally potent. A more realistic increase in fN2O of $10-1$ times would have provided α not enough the Protein war added to the effect of CH4, the net warming is about 10!. mostly on the magnitude of its source flux. But this is not the can he veralued tun be resolven about 230 nm (Kaiser et al., 2003). Those wavelengths are

blocked by O2 (in the Herzberg bands), so as atmospheric O2

dance for two different CH4 concentrations, 1.6 ppmv (A) and 100 ppmv (B).

USNO Scientific Colloquium, 01/26/17 mates. A CO2 mixing ratio of !0.2 and !0.01 would have sufficed in the early and late $\mathcal{S}/17$ if the early and late $\mathcal{S}/17$ if the early if the early if the early $\mathcal{S}/17$ \mathbf{F} is \mathbf{F} as a function of \mathbf{F} as a function of \mathbf{F}

> In combination with 1 ppmv N2O and 100 ppmv CH4, however, a CO2 mixing ratio of only !0.06 would have

> dance. The two solid curves represent climate model calculations for two solar luminosities: 83% and 94% of present value So. The N2O and CH4 concentra-

Young Sun

Hadean Earth

 The Pathway to Bio Molecules

Protons are 10⁶ times more efficient!

What organic molecules do if given energy + Darwinian evolution Life is a self sustaining chemical system capable of evolution

Energy flux

Life is a chemical soliton! Life is a chemical soliton! **The ascent of molecules** *Proton irradiation forms amino acids, nucleobases* $HCN + H₂O$ --> $NH₂CHO$ -formamide

> In 2010, while pondering the question of why chemistry ofen gets a rough ride in the mainstream media, this journal suggested1

> Te classic experiment relationship to the chemical origins of life, describing the formation of amino acids when an electric discharge was passed through mixtures of methane, ammonia, water and hydrogen \mathcal{L} with \mathcal{L} as the Miller–Urey experiment — was

 by Stanley Miller in *Science* in 1953. So, although this is not a new area of research, it is considered to the hast of the hast seen a recent recent results in the centre of interest4 diferent aspects of chemistry5

together in this interdisciplinary feld that encompasses topics such as the prebiotic synthesis of nucleosides and nucleotides, the

this occurred before or afer the formation

Life Cookies istema in portance of the systems of students in the system in the system in the system in the system in the s understanding the prebiotic world. In their

on the early Earth could have evolved into the complex dynamic biochemistry that we know today.

a prebiotically plausible synthesis of the

Tese articles describing and discussing research on this topic are complemented by an interview with Matthew Powner, now based at University College London. Powner worked with both the Sutherland and Szostak research groups, as a PhD student and a postdoctoral researcher, respectively, and is a co-author of both Articles that appear in this issue. In the interview, Powner describes how he sees the two pieces of work ftting together as part of "a colossal chemical jigsaw", and also discusses his own nascent independent research career.

Powner's passion for this subject is evident from the answers, not least when he gives reasons for why this topic should be funded. We asked him this question because, in these straightened economic times, a case ofen needs to be made (particularly to funders) for why a given piece of research demonstrates an obvious long-term societal (usually monetary) beneft. Powner points to our fundamental thirst for knowledge and compares the funding of this feld to that of the Large Hadron Collider. Te unanticipated benefts of such blue-sky research are frequently cited as the best reason for it to be supported. It seems reasonable to suggest that such benefts may be likely to arise from interdisciplinary $t_{\rm{max}}$ as the origin of life, because understanding them requires the collaboration of researchers from so many diferent specialties. In attempting to answer questions as fundamental as "where do we

latest work, the Sutherland group show that similar systems-chemistry approaches can explain the important region of problem. A combination of problem. plausible reactions and site-selections and reaction that results in acylation of the $20-40$ and the promotes promotes promotes promotes promotes and thus promotes p ligation in the desired 3ʹ,5ʹ-fashion.

Te other Article, from Jack Szostak and co-workers, takes a complementary

the backbone of polynucleotides irregular,

the existence — of enzymes.

. By examining the activity of RNA molecules containing a mixture of 2ʹ,5ʹ and 3ʹ,5ʹ linkages, they show that a mixture of isomers may not have been a problem for the evolution of RNA. Indeed, they

 \mathcal{A}

Making "Cookies" of Life on Earth/Mars

What Is Habitability ? Not life, but physico-chemical conditions supporting life As We Know It

- *Rocky planet with a surface liquid water (CHZ)* • *Plate tectonics with suduction* • *Presence of a planetary magnetic field* • *Stability of complex carbon molecules* • *CHNOPS and a range of metals (Fe, Mn) to build molecules and act as catalysts*
- *Presence of energy sources to drive bio processes and evolution or redox reactions in the crust*

" *Long life expectancy (> 1 Gyr)* " *Be luminous enough - the planet doesn't have to be too close How To Become a "Star" Parent ?*

• *stars with M < 0.5M_o (M) will tidally lock* • *XEUV/wind flux and atmospheric loss* " *Have high "metallicity*" " *Special constraints on a binary system* " *"Child" planet needs an emotionally stable parent*

 Mid aged G-K stars are the best "parents"

Strategies to Search for Life

Earth analogs within BZ **R=2,700/NIRSpec**

 Direct Imaging Young Earths "pregnant" with life N2O- 2.9, 4.5, 7.9 µm, HCN – 3 and 14.3µm Beacons of Life Earth twins: NO – 5.3 µm at 100 TW OH at 100 TW at 1.6 and 2 µm

• *Extension of FKSI* • *Four 2.5 m telescopes* • *Boom ~ 25 m* • *IWA~20 (/5) mas*

Exo Life Beacon Space Telescope

