

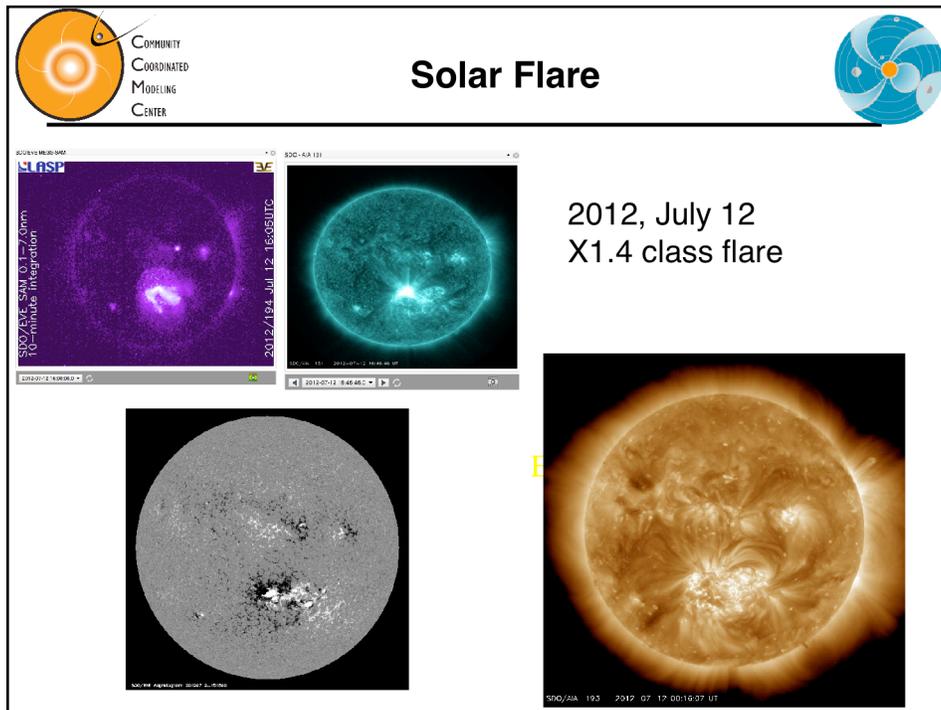
Flares, CMEs, Solar Wind High Speed Streams

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In this presentation I will be talking about solar flares and coronal mass ejections and their relation to the space weather



Solar flare is a sudden brightening observed over the Sun's surface or the solar limb, which is interpreted as a large energy release. Flares are mainly followed by a mass ejections from the solar atmosphere called coronal mass ejections. The flare ejects clouds of electrons, ions, and atoms through the corona of the sun into space.

In this slide is shown the solar flare of 2012 July 12.

The three color images show SDO satellite (<http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Glossary>) images of the same flare in three different wave length intervals: 1) the left is SDO EVE X-ray image (http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Full_iSWA_Cygnets_List

The central panel shows SDO AIA extreme ultraviolet image for 13.1 nm http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/SDO_AIA_131

and the right hand side movie - for a longer length of 19.3 nm http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/SDO_AIA_193

In this slide is shown also in the lower left corner a magnetogram (<http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Glossary>) - magnetic field map - of the solar photosphere for the same period of time. This flare is associated to the large and magnetically complicated active region in the southern hemisphere of the solar surface as is usually takes place.

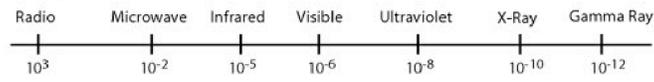


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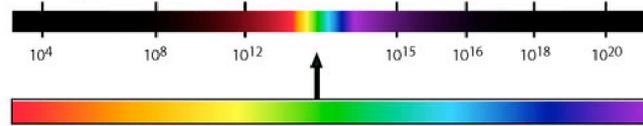
A powerful solar flare radiates across the whole electromagnetic spectrum

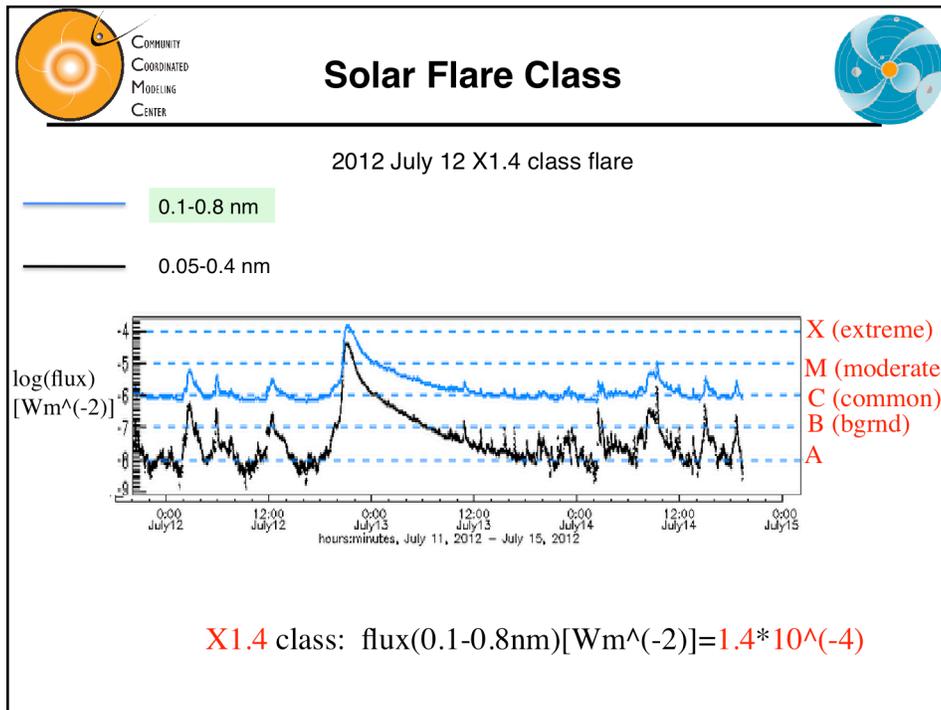
THE ELECTRO MAGNETIC SPECTRUM

Wavelength
(metres)



Frequency
(Hz)





In this slide is shown the time line for the X-ray flux intensity

(http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/GOES_X-RAY) for the same flare of 2012 July 12, in two different wave length intervals . A strong flare usually manifests itself by a sudden a jump of the X-ray intensity.

The flares are characterized by classes according to their X-ray flux intensity in the 0.1 – 0.8 nm wavelength range.

Flux[Wm⁽⁻²⁾] > 10⁽⁻⁴⁾ X - (eXtreme), strong

Flux[Wm⁽⁻²⁾] > 10⁽⁻⁵⁾ M – (Moderate)

Flux[Wm⁽⁻²⁾] > 10⁽⁻⁶⁾ C – (Common)

Flux[Wm⁽⁻²⁾] > 10⁽⁻⁷⁾ B – (Background)

Flux[Wm⁽⁻²⁾] > 10⁽⁻⁸⁾ A

M1.4 class flare that at the peak the flare intensity was equal to:

$$\text{flux/Wm}^{(-2)}=1.4*10^{(-4)}$$



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Quick quiz



Do you think that X-class flares can happen close to the solar minimum?

Solar material is in constant, active motion. Plasma motion in the convective zone generates the magnetic field which plays a crucial role in the physical processes on the Sun and its activity.

In this slide is shown solar atmosphere relatively close to the photosphere seen by the Japanese HINODE satellite. You can notice constantly “boiling” surface, together with numerous small eruptions, brightening, fountains of the solar material. In the left corner solar plasma is flowing along the arcade shaped magnetic field line loops (<http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Glossary>). This flow happens because, although charged particles experience difficulty in moving across, they can almost freely move along the magnetic field lines.

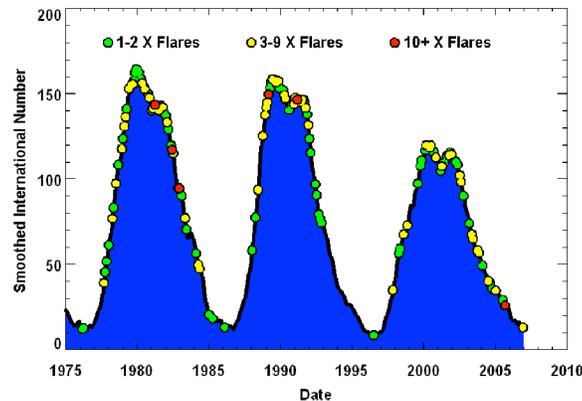


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Flares over the Solar cycle



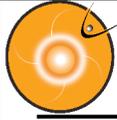
Solar flares have been monitored by x-ray detectors on GOES satellites since 1976. The number of X-Class flares per month increases with the number of sunspots but **big flares can occur anytime sunspots are present.**



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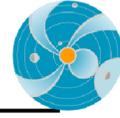
In this plot the frequency of the X-class flares is plotted over the sunspot number plot. Green dots represent 1-2 X-class flares, yellow – 3-9 X-class flares and red – 10+ X class flares.

You can see that the number of X-Class flares per month increases with the number of sunspots, but big flares can occur anytime sunspots are present.



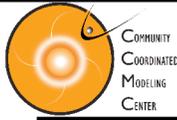
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Flare Characteristics



- Flares tend to occur in isolation, localized in space and time but with strong correlations.
- Typically one active region will produce dozens of flares, especially during periods of flux emergence (often near the beginning of the lifetime of a given region, but not always).
- The most powerful events usually occur in active regions.

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Flare: SWx impacts



- Cause radio blackout through changing the structures/composition of the ionosphere (sudden ionospheric disturbances) – x ray and EUV emissions, **lasting minutes to hours**
- Affect radio communications, GPS, directly by its radio noises at different wavelengths
- Contribute to SEP – proton radiation, **lasting a couple of days**

8

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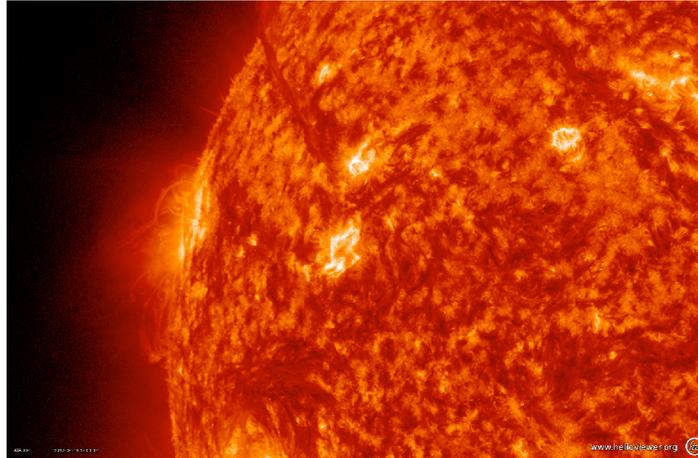
Affect radio communications, GPS, directly by its radio noises at different wavelengths

Contribute to Solar Energetic Particle events (SEP) – proton radiation, **lasting a couple of days**



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Coronal Mass Ejection



Solar flares are often accompanied by so called Coronal Mass Ejection (CME), when huge mass of solar plasma erupts from the solar atmosphere. <http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Glossary>).

If CME is directed towards Earth, it can reach the Earth in 1-3 days, depending on its speed and size. CMEs are usually causing the strongest geomagnetic storms.



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Quick quiz



How do you think CMEs
are correlated with
solar flares?

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In this slide is shown solar atmosphere relatively close to the photosphere seen by the Japanese HINODE satellite. You can notice constantly “boiling” surface, together with numerous small eruptions, brightening, fountains of the solar material. In the left corner solar plasma is flowing along the arcade shaped magnetic field line loops (<http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Glossary>). This flow happens because, although charged particles experience difficulty in moving across, they can almost freely move along the magnetic field lines.

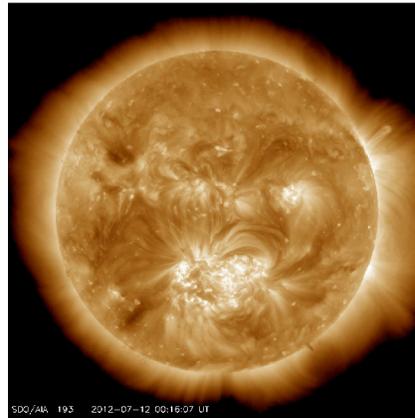
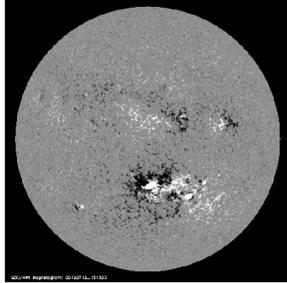


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CME/flare

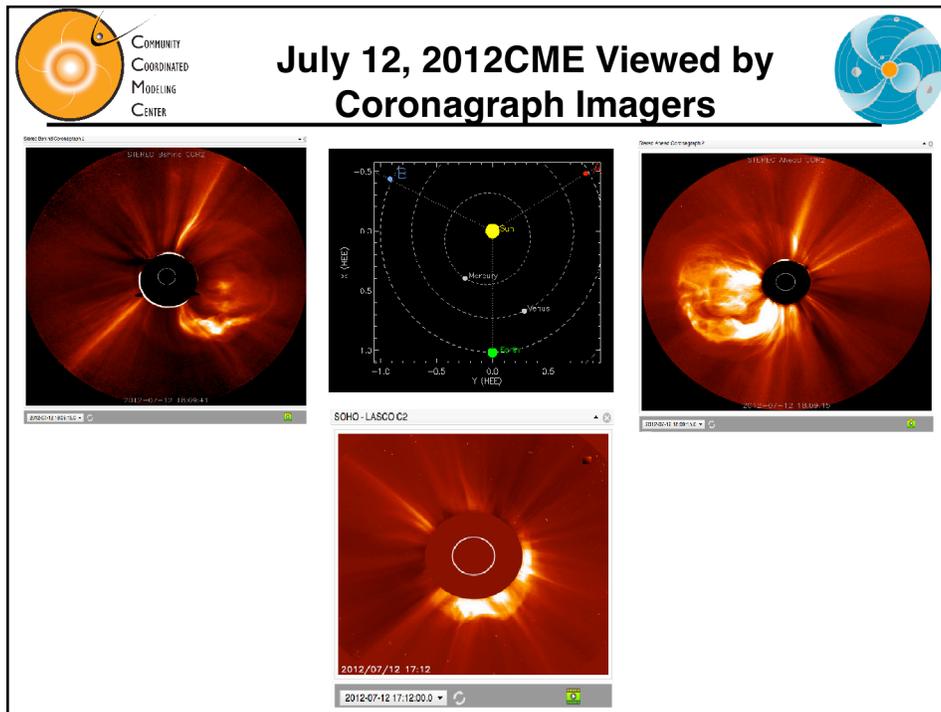


The most energetic CMEs occur in close association with powerful flares in the active regions. Nevertheless large-scale CMEs do occur in the absence of major flares even though these tend to be slower and less energetic.



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When strong flare/CME occurs, it gives off emission across the whole electromagnetic spectrum, at the same time energetic particles.



But the better view of a CME and its motion in the interplanetary space can be obtained from the coronagraph images. Coronagraphs create an artificial eclipse of the sun. Eclipses allow corona to be better viewed, but natural eclipses do not happen often. Occulting disk blocks the bright sun so we can observe corona features.

This slide shows the same CME seen in 3 coronagraphs located on three different satellites. The schematic in the upper center shows the location of STEREO B, SOHO and STEREO A satellites in the ecliptic plane at the time when the CME occurred.

The CME was moving sort of towards the Earth, so STEREO B sees it moving to the right in its plane of sky, for STEREO A it moves to the left and it's a halo image for SOHO that is located on a sun-earth line.

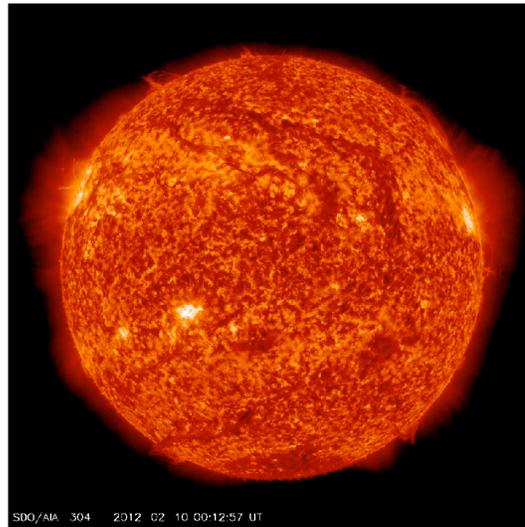


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CME from a Filament Eruption



Northeast (upper left) quadrant starting around 19:00 UT on Feb 10, 2012



SDO/AIA 304 2012_02_10_00:12:57 UT

Not all CMEs originate from the active regions. Here is shown an example of a CME caused by so called filament eruption. Filaments are formed in magnetic loops that hold relatively cool, dense gas suspended above the surface of the Sun. Most of the filaments are created by stretching of the active regions with two different polarities due to the differential motion of the solar photosphere material. Magnetic instabilities cause filaments to raise and parts of it disconnect from the solar surface resulting into CMEs.



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CME Properties



- Mass: $\sim 10^{15-16}$ g
- Speed: few hundred - 3000km/s

..or

- Mass: ~ 1 million Nimitz-class aircraft carriers
- Speed: 1.5 -10 million km/hour



Earth?



- Arrives to Earth in 1-2 days

CME properties:

- Mass: $\sim 10^{14}$ kg
- Speed: few hundred - 3000km/s

In other words:

- Mass: ~ 1 million Nimitz-class aircraft carriers
- Speed: 1.5 -10 million km/hour

If it hits the Earth, it can have a significant impact.



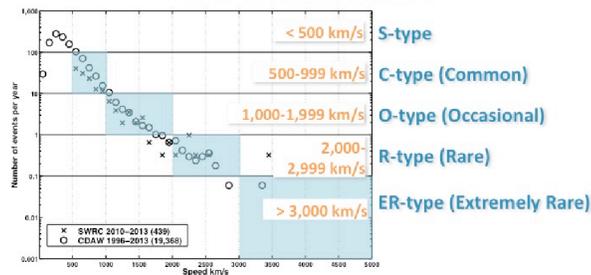
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CME SCORE



- A new category system for CMEs based on frequency of detection and speed
- Complements Flare Classes
- Applicable in space weather operations and research

Space Weather Research Center CME SCORE Scale

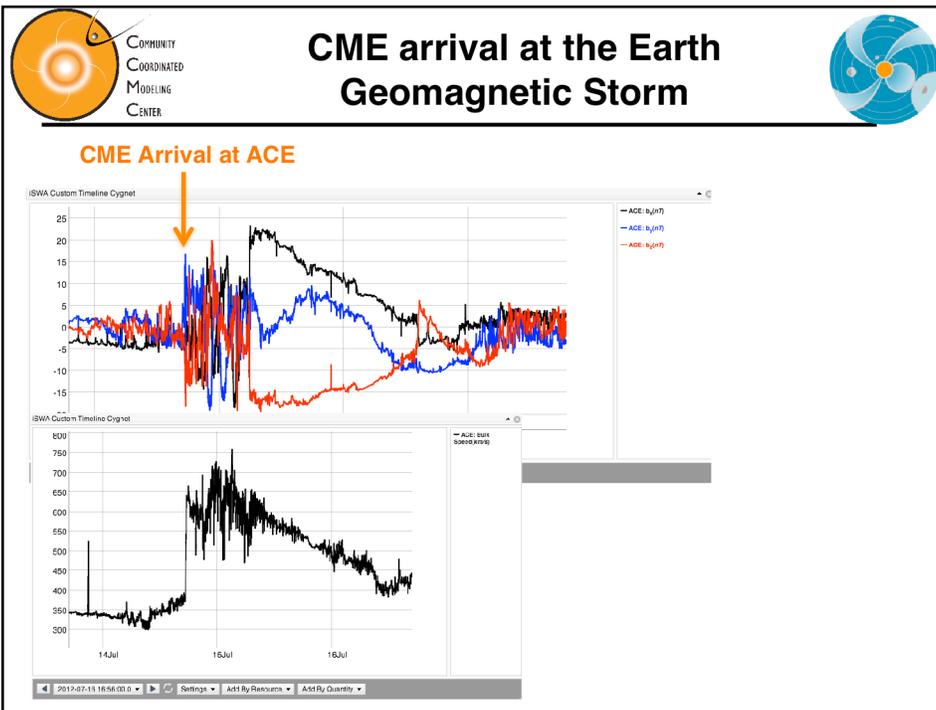


Recently we, here at SWRC, introduced, as we call it, CME SCORE scale, a simple new category system for CMEs based on frequency of detection and speed.

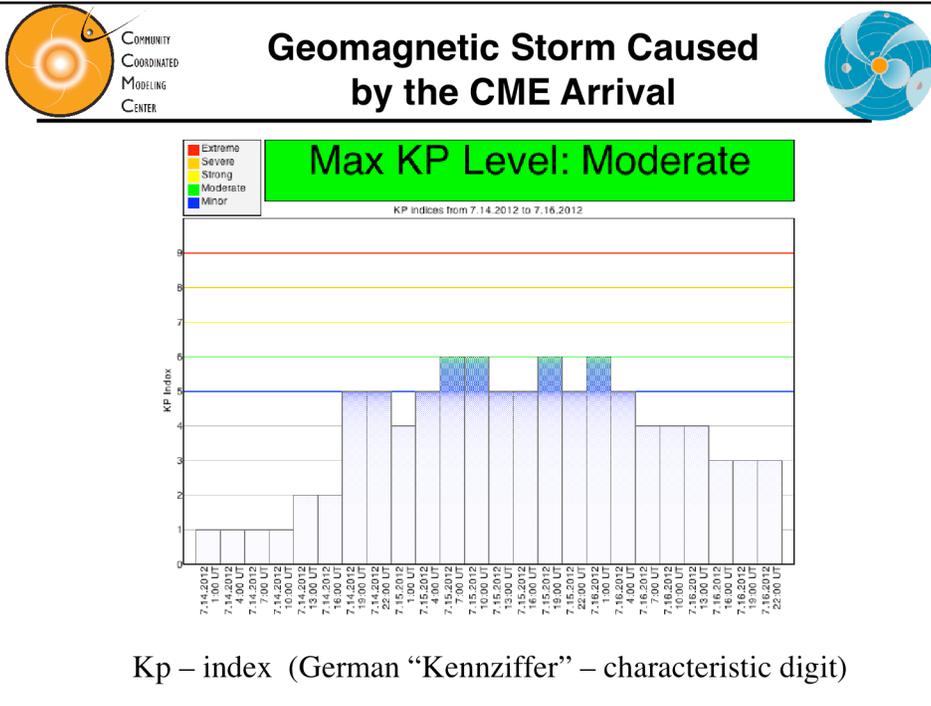
It complements flare classes and is applicable in space weather operations and research.

The graph here shows the frequency of the CME detection (number of events per year) as a function of the CME speed. Bin size is 100 km/s.

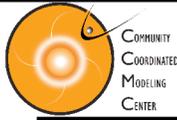
So we have S-type CMEs with the speed less than 500/km/s, C-type (common) for the range 500-999, O-type (occasional) for 1000-1999, R-type (rare) for 2000-2999 and ER-type (extremely rare) CMEs for CMEs with the speed more than 3000 km/s.



When the CME arrived to the ACE satellite, close to the Earth, it caused a shock seen here in the magnetic field and the solar wind speed measurements by ACE. This impact caused a strong disturbance in the Earth's magnetosphere or as they call it geomagnetic storm.



The strength of the geomagnetic storm is measured by so called Kp index, which characterizes degree of the disturbance of the Earth’ magnetic field ground base measurements. Kp goes from 0 to 9. This storm with Kp=6 is considered to be a moderate one.



Quick quiz



What CME parameters do you think are important for the strength of the geomagnetic storm?

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SWx impacts of CME



- Contribute to SEP (particle radiation): 20-30 minutes from the occurrence of the CME/flare
- Result in a geomagnetic storm: takes 1-2 days arriving at Earth
- Result in electron radiation enhancement in the near-Earth space: takes 1-3 days

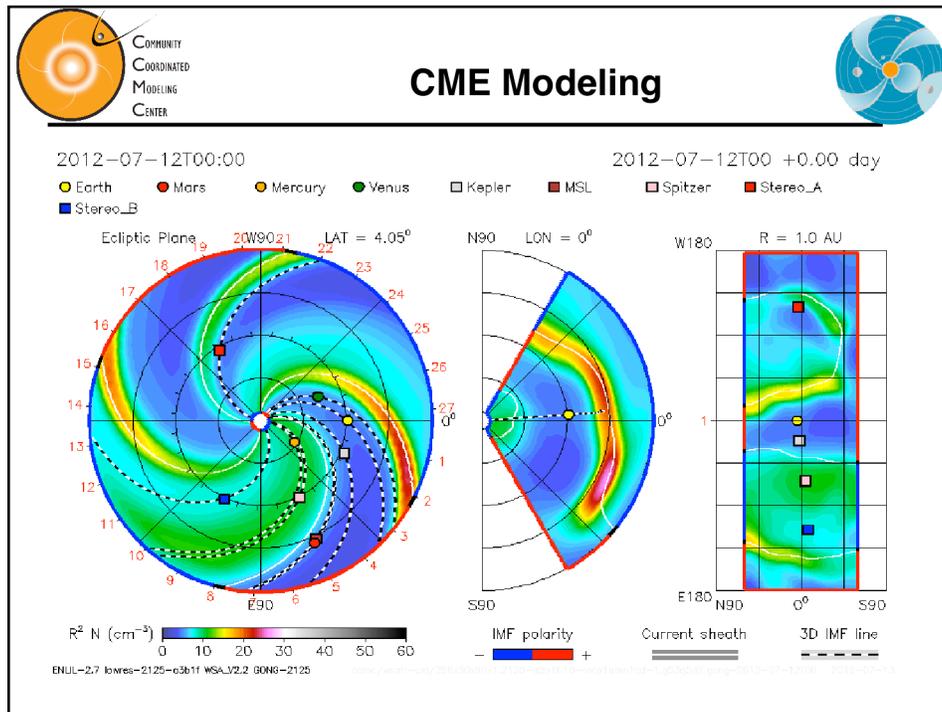
Affecting spacecraft electronics – surfacing charging/
internal charging:
Radio communication, navigation, power grids, pipelines,
and so on

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Contribute to SEP (particle radiation): 20-30 minutes from the occurrence of the CME/flare

Result in a geomagnetic storm: takes 1-2 days arriving at Earth

Result in electron radiation enhancement in the near-Earth space: takes 1-3 days.



This movie shows WSA-ENLIL modeling of the same July 12, 2012 CME. In all three panels normalized solar wind density is plotted.

In the left panel is shown the view from the Solar north pole down to the ecliptic plane, where all the planets are moving around the sun. Yellow circle is the Earth, orange, closest to the sun, is the Mercury, green – Venus, red – Mars, red and blue squares are STEREO A and B satellites respectively. Strictly speaking the first panel shows cut through a plane passing through the Earth and is parallel to the Sun's equatorial plane.

The second plane is a meridional cut passing through the Earth. The latitude goes from -60 to 60 degrees.

The third panel shows Longitude-Latitude map of a quasi-sphere at 1 AU cut for latitudes with absolute values more than 60 deg.

You can see that CME causes enormous disturbance throughout the whole heliosphere. Sometimes CMEs are called space hurricanes. For terrestrial (regular) weather we need to know the propagation path and strength of a hurricane. The same way in space weather we need to know the propagation path and strength (speed, mass) of a CME to estimate its possible impact.