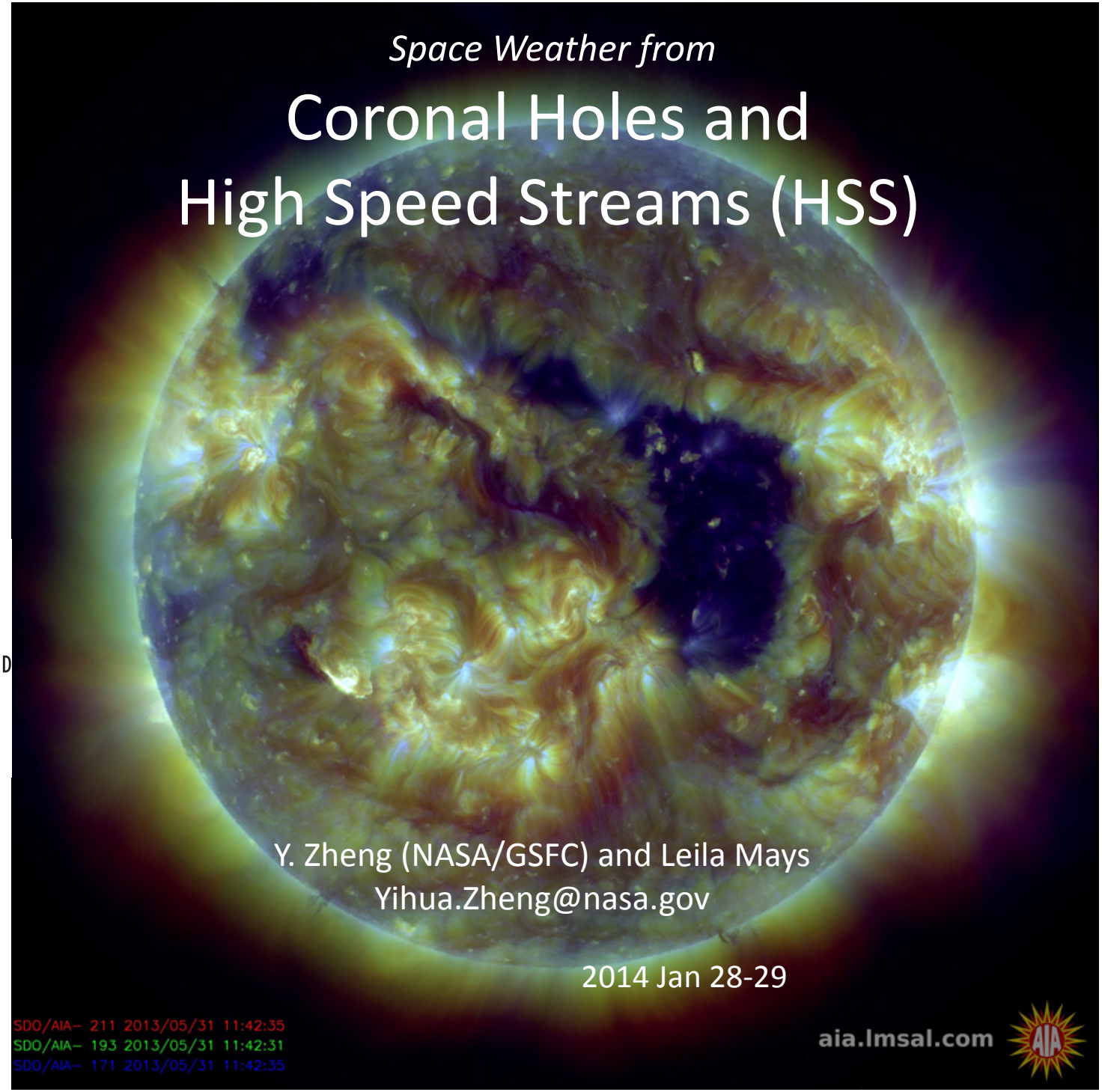

Space Weather from
Coronal Holes and
High Speed Streams (HSS)
Y. Zheng (NASA/GSFC) and Leila Mays
Yihua.Zheng@nasa.gov
2014 Jan 28-29
SDO/AIA– 211 2013/05/31 11:42:35
SDO/AIA– 193 2013/05/31 11:42:31
SDO/AIA– 171 2013/05/31 11:42:35
aia.lmsal.com

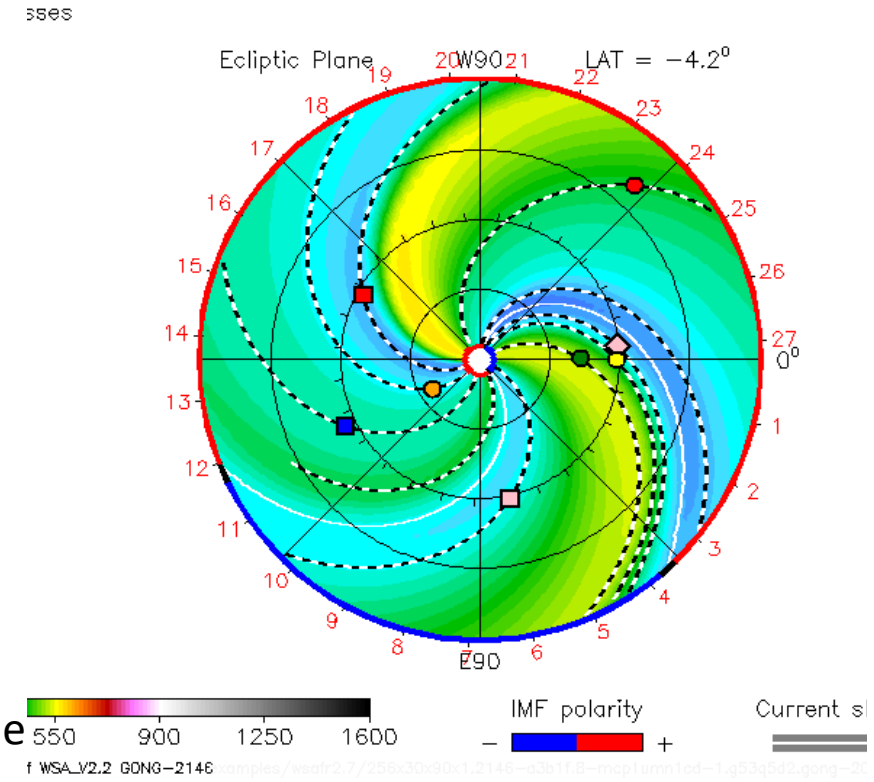
Solar Wind

The **solar wind** is a flow of **plasma** and the frozen-in solar **magnetic field** from the Sun. The outward flow is due to the gas pressure difference between interplanetary space and the solar corona.

Changes in the solar magnetic field (from solar activity) influence the solar wind which, in turn, influences planets, spacecraft, and other bodies inside the solar wind (the **heliosphere**).

TABLE 4.1. Observed Properties of the Solar Wind near the Orbit of the Earth (1 AU)

Proton density	6.6 cm^{-3}
Electron density	7.1 cm^{-3}
He^{2+} density	0.25 cm^{-3}
Flow speed (nearly radial)	450 $km \cdot s^{-1}$
Proton temperature	$1.2 \times 10^5 K$
Electron temperature	$1.4 \times 10^5 K$
Magnetic field (induction)	7×10^{-9} tesla (T)



near the Earth, compared to the magnetosphere, the solar wind plasma (mostly ionized Hydrogen) is hot, tenuous, and fast moving, and the weak magnetic field is nearly parallel to the ecliptic plane, but 45˚ to the Sun-Earth line

The **Parker spiral**, is the spiral of Archimedes magnetic geometry of the solar wind due to solar **rotation**. Parcels of solar wind leaving the sun are analogous the water spirals formed from a rotating sprinkler. The angle a solar wind magnetic field line makes at 1 AU is close to 45 degrees.

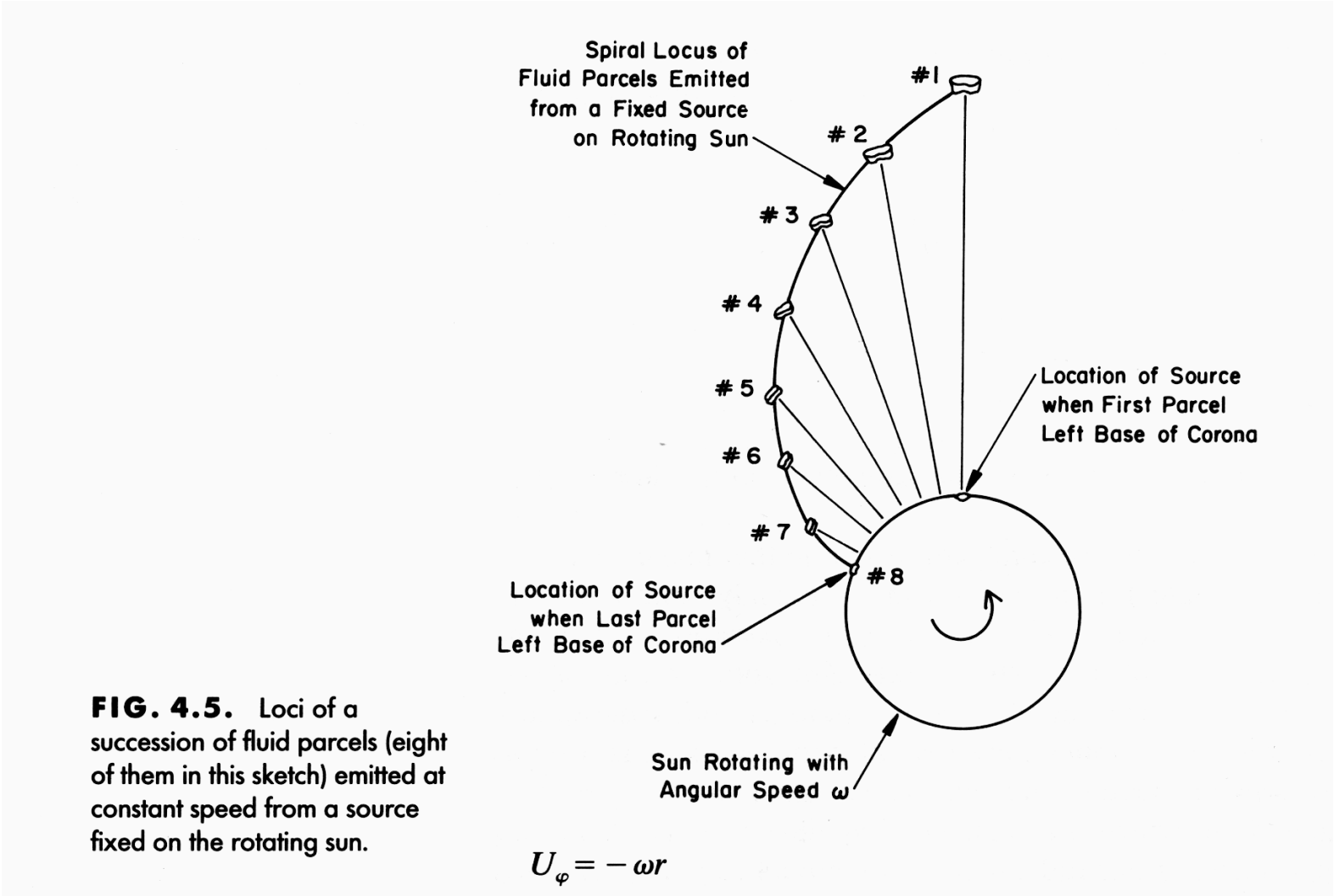
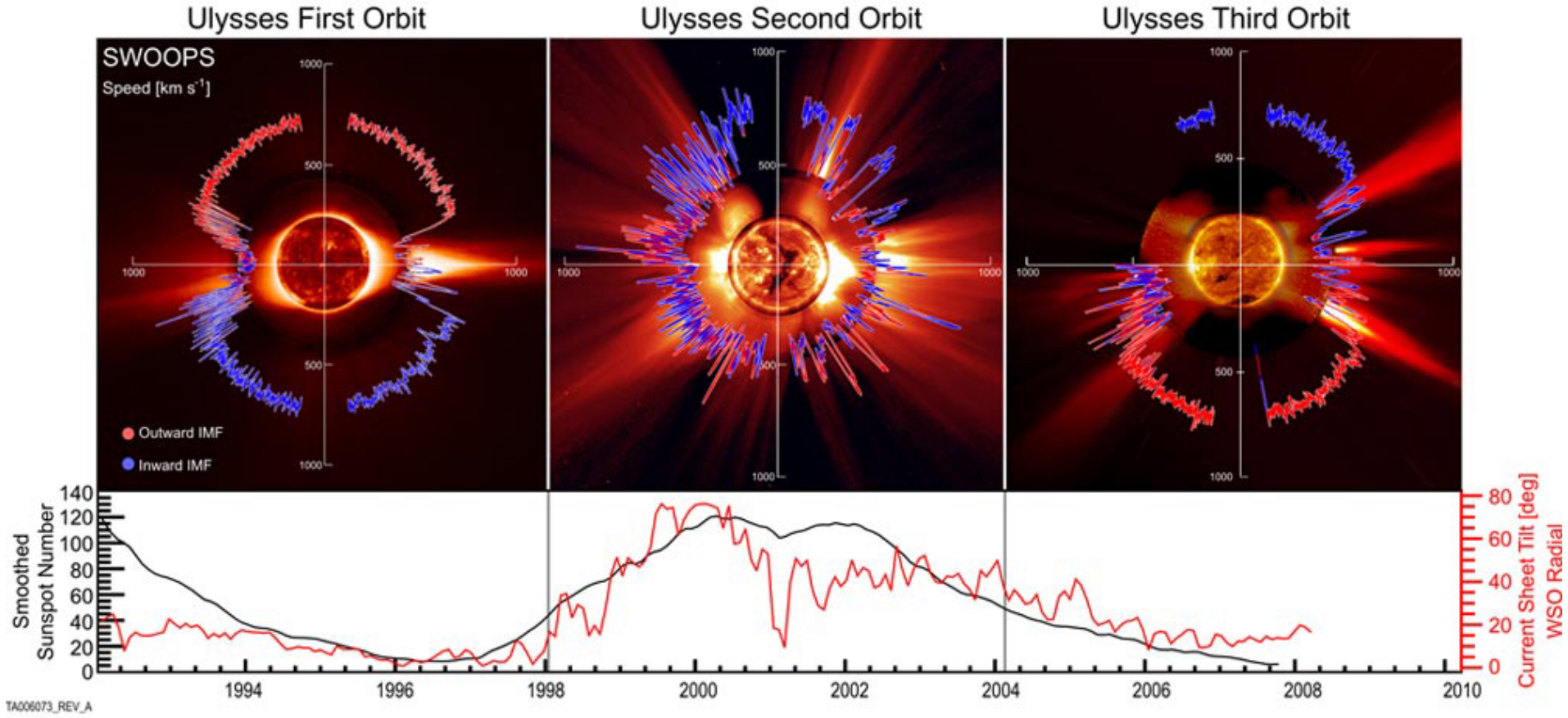


FIG. 4.5. Loci of a succession of fluid parcels (eight of them in this sketch) emitted at constant speed from a source fixed on the rotating sun.

Solar wind can be divided into fast and slow wind components.



solar minimum

solar maximum

solar minimum

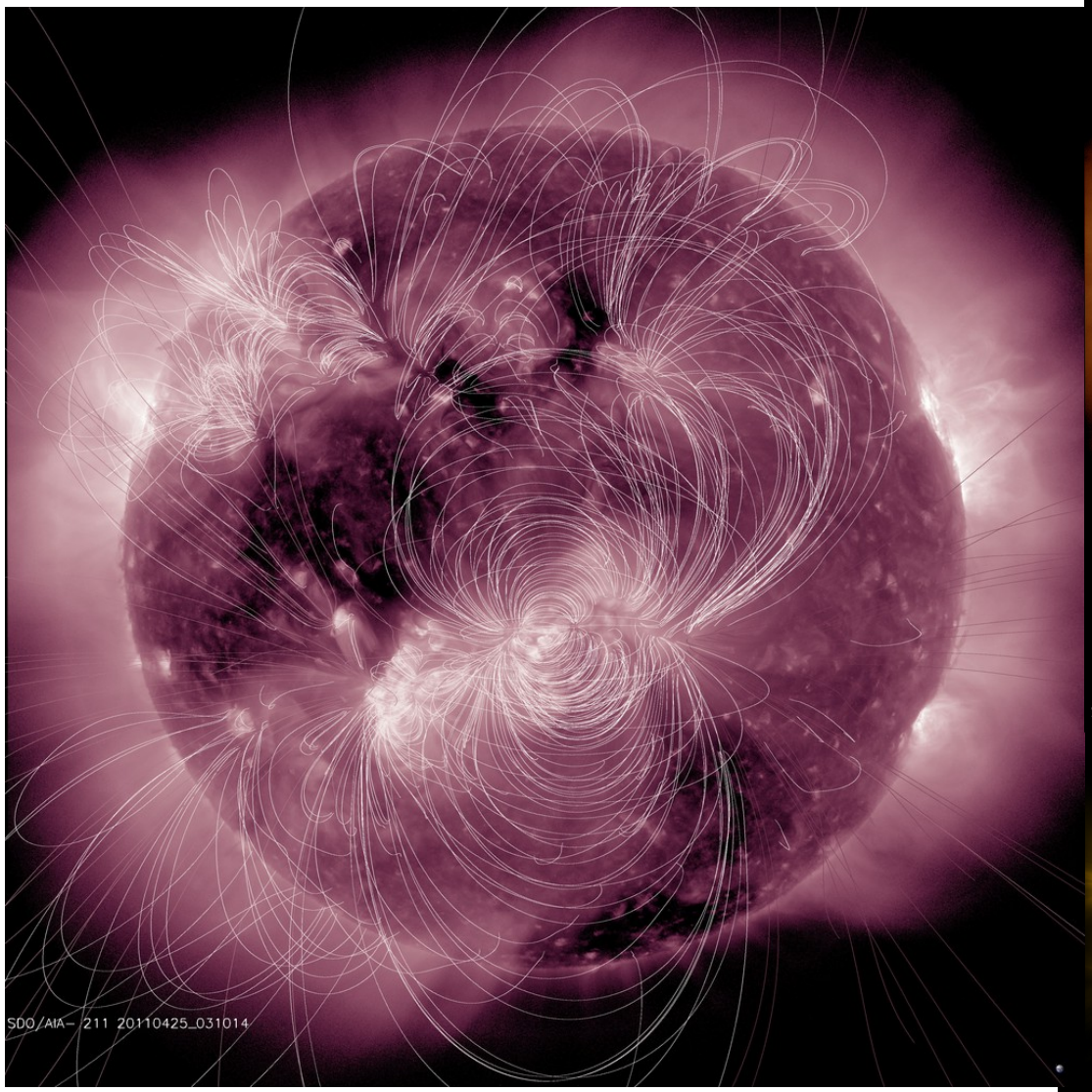
Solar wind can be divided into fast and slow wind components.

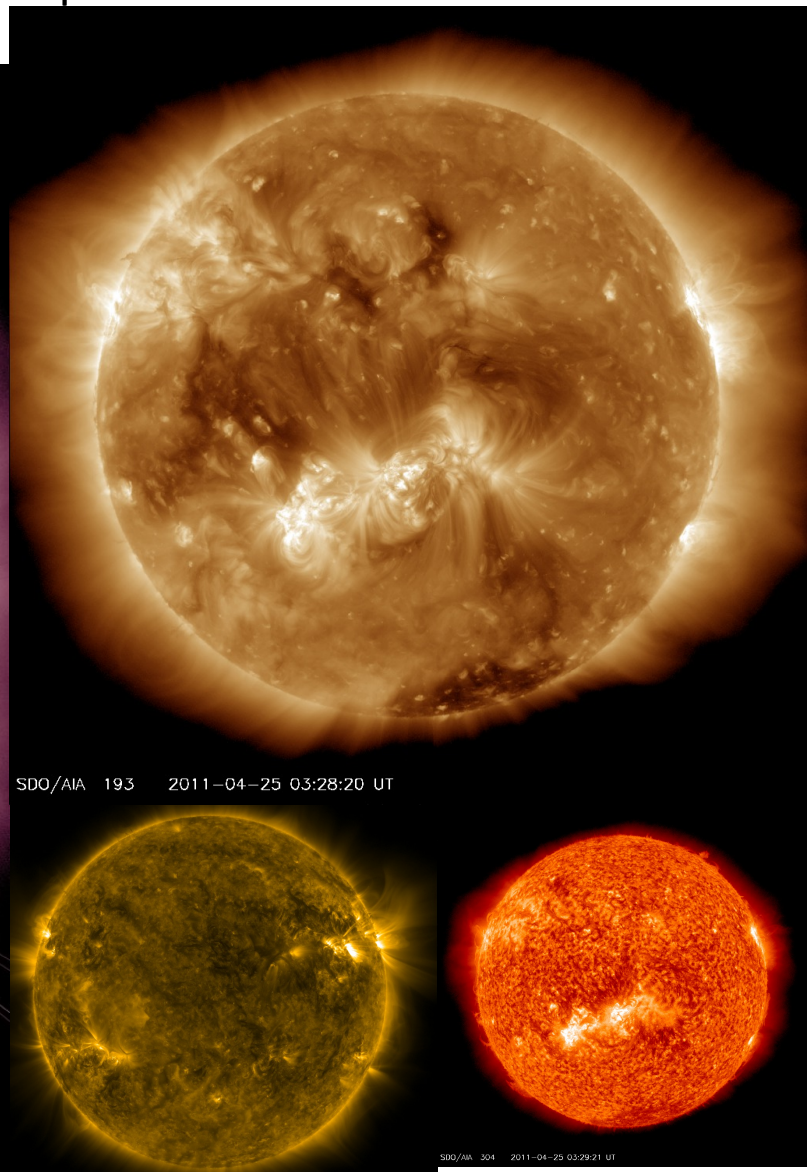
Fast wind	Slow wind
450–800 km/s	<~450 km/s
$n_P \sim 3\ cm^{-3}$	$n_P \sim 7–10\ cm^{-3}$
~95% H, 5% He, minor ions and same number of electrons	~94% H, ~4% He, minor ions and same number of electrons – great variability
$T_P \sim 2 \times 10^5$ K	$T_P \sim 4 \times 10^4$ K
$B \sim 5$ nT	$B \sim 4$ nT
Alfvénic fluctuations	Density fluctuations
Origin in coronal holes	Origin 'above' coronal streamers and through small-scale transients

(Bothmer and Zhukov, 2007)

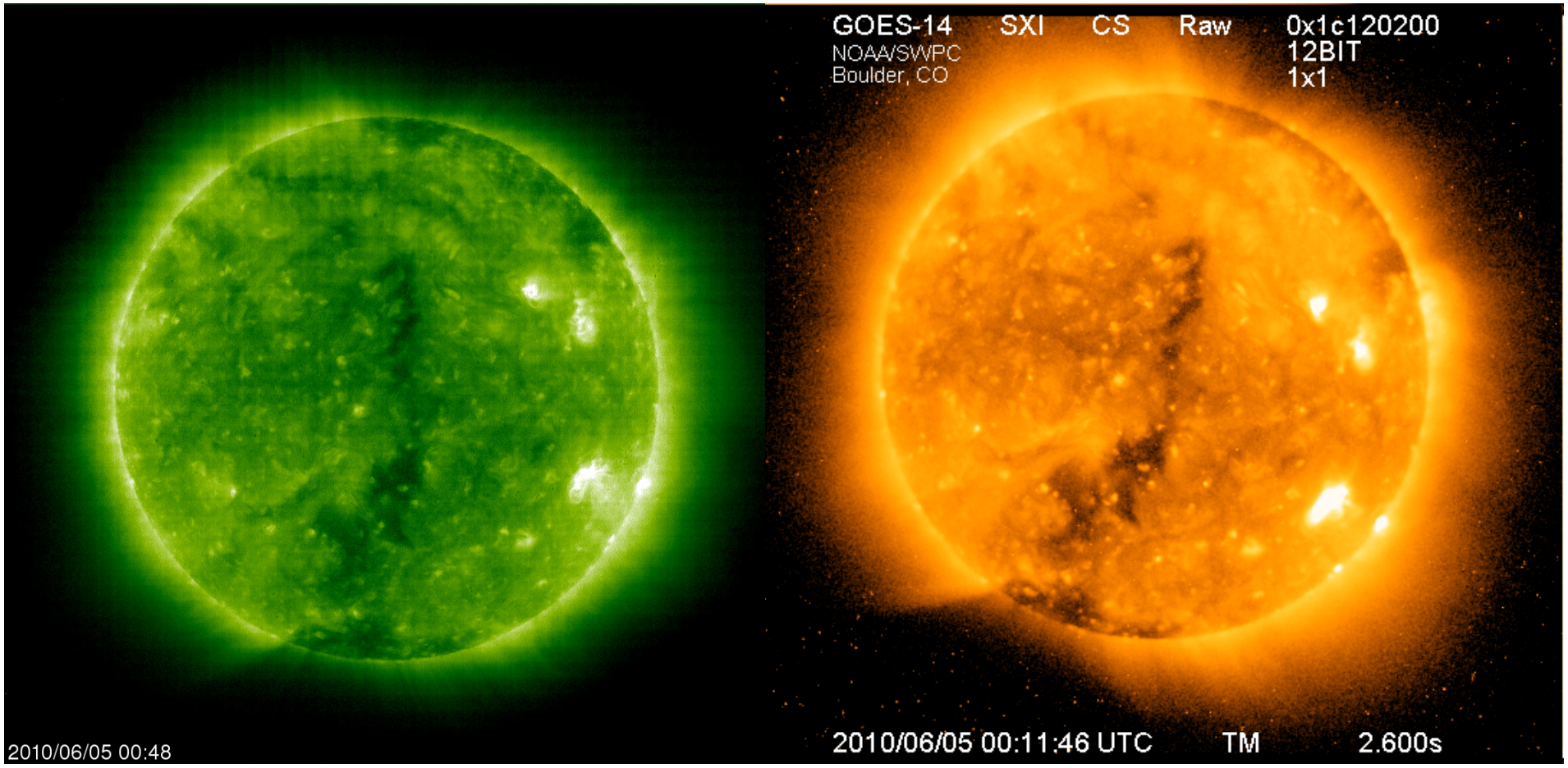
Coronal Hole High Speed Streams

Coronal holes appear as dark areas on the solar surface in the **EUV** (extreme ultraviolet) and **X-ray** radiation. They have a **lower density** and **temperature** compared with the surrounding corona. **Coronal holes** correspond to regions of open magnetic fields. Visible best in lines with temperatures more than 10^5 K.





Large **polar coronal holes** are persistent for about 7 years around solar minimum. During solar maximum and high solar activity they exist at all latitudes, but are less persistent.



Coronal holes correspond to regions of open magnetic fields.

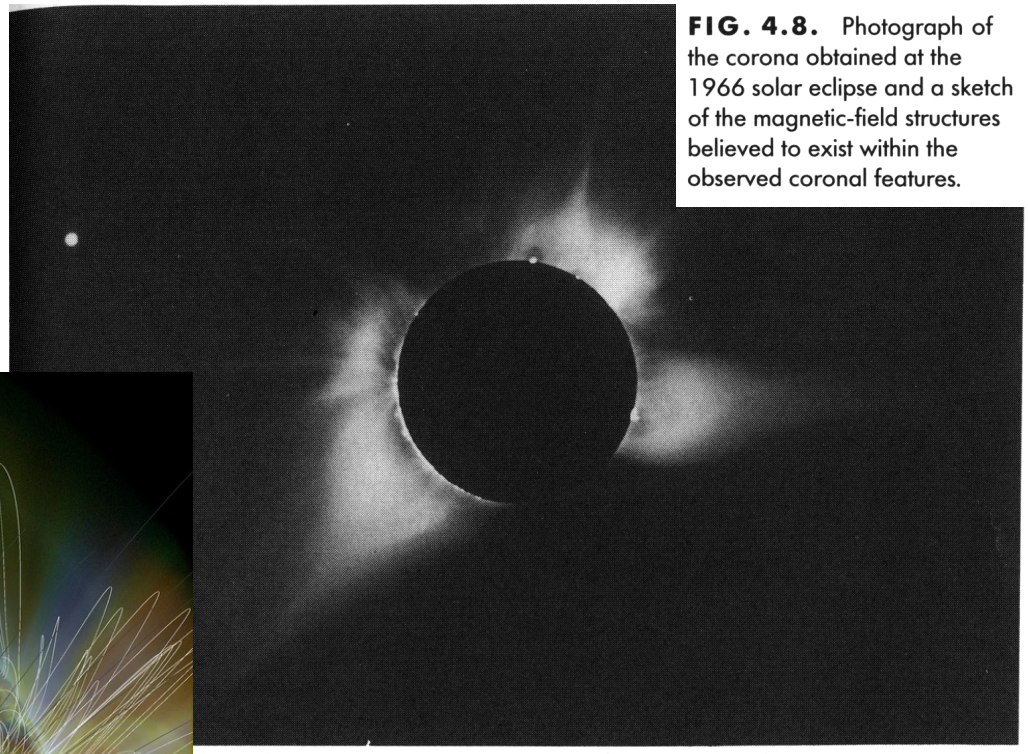
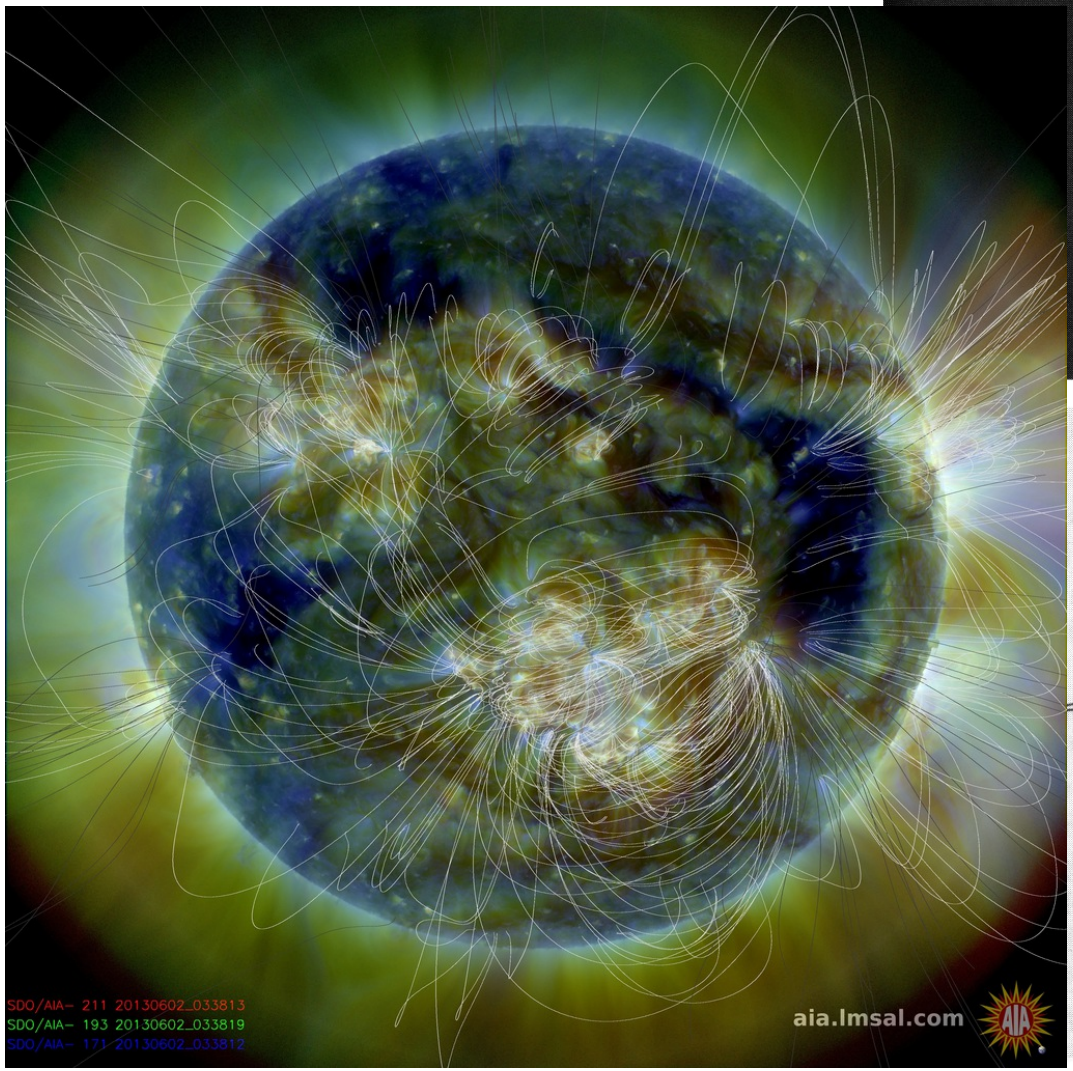
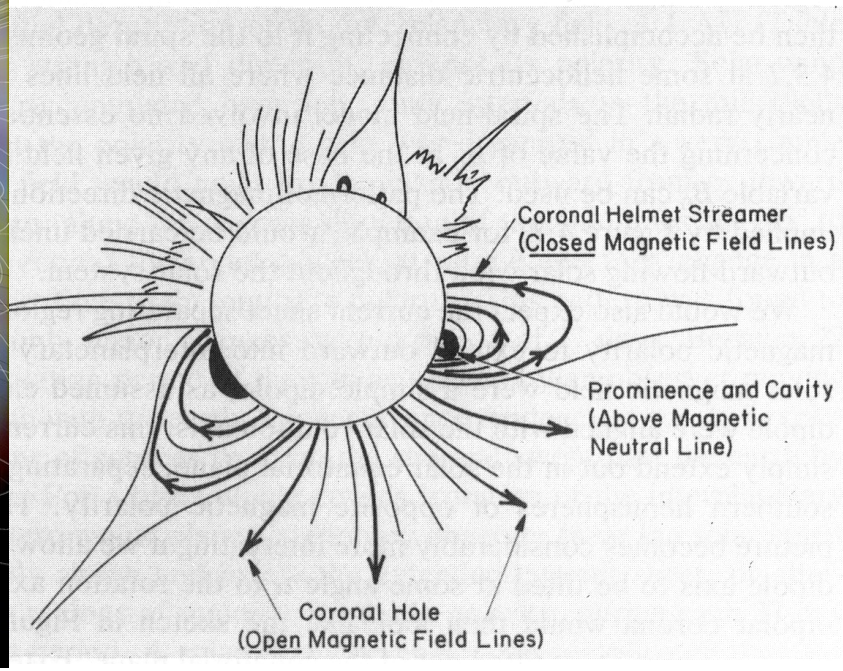


FIG. 4.8. Photograph of the corona obtained at the 1966 solar eclipse and a sketch of the magnetic-field structures believed to exist within the observed coronal features.





High speed solar wind streams are formed by higher speed solar wind originating from **coronal holes**. Higher speed streams are less tightly wound in the Parker spiral compared to slower ones, and at various distances the faster solar wind **overtakes** the slower wind ahead of it.

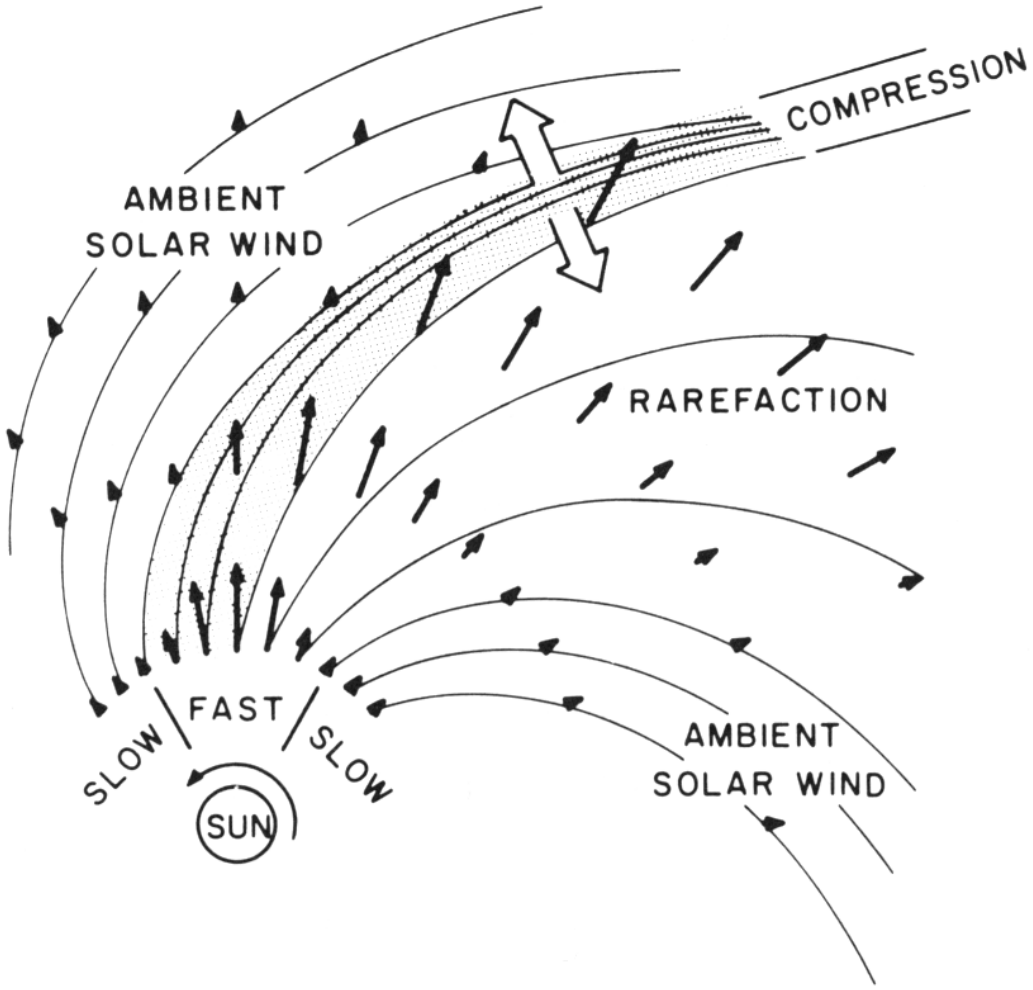


FIG. 4.13. Geometry of the interaction between fast solar wind (on less tightly wound spiral streamlines) and ambient solar wind (on more tightly wound spiral streamlines). The plasma is compressed where streamlines converge. (From Pizzo, 1985.)

A **stream interaction region** (SIR) forms at the compressed boundary between the fast and slow solar wind in a high speed stream. High speed streams from persistent coronal holes over multiple solar rotations are called **corotating interaction regions** (CIRs).

CIR is a subset of SIR

More long-lasting

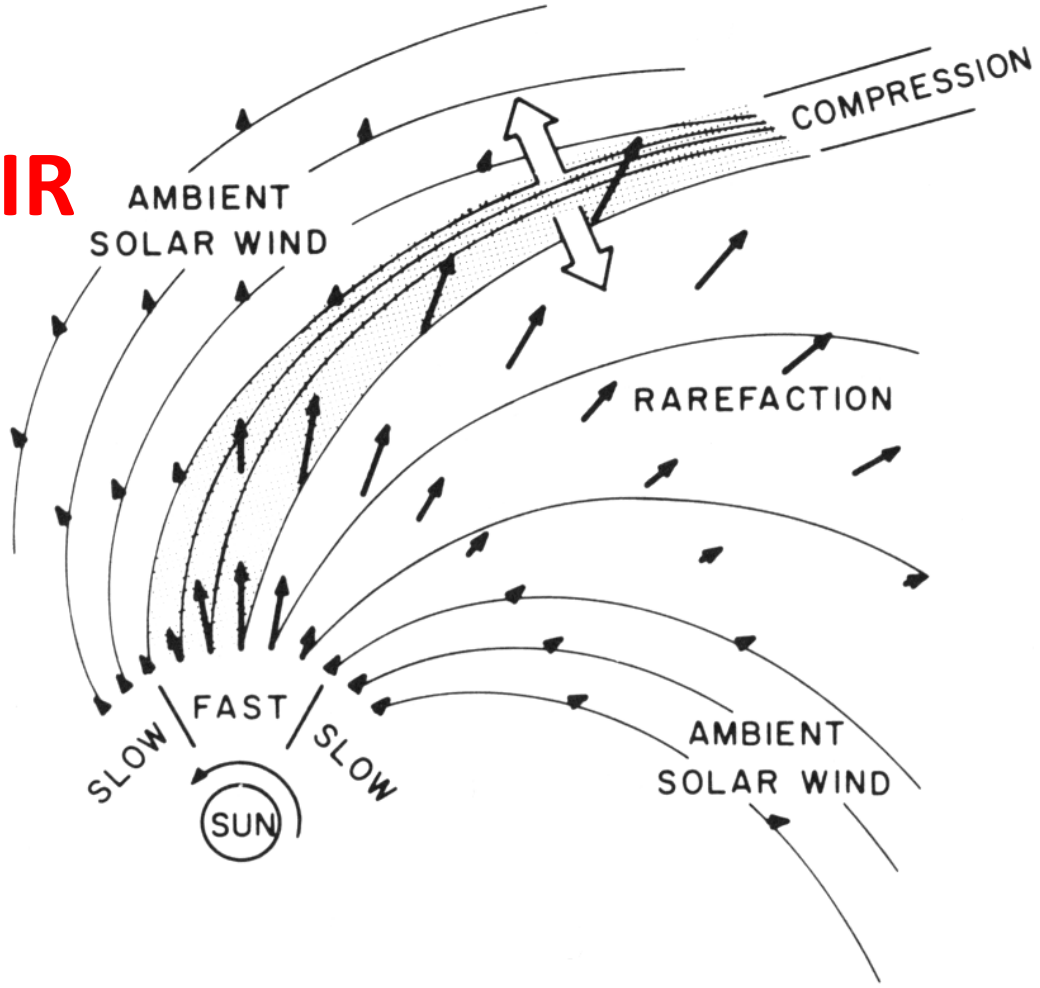
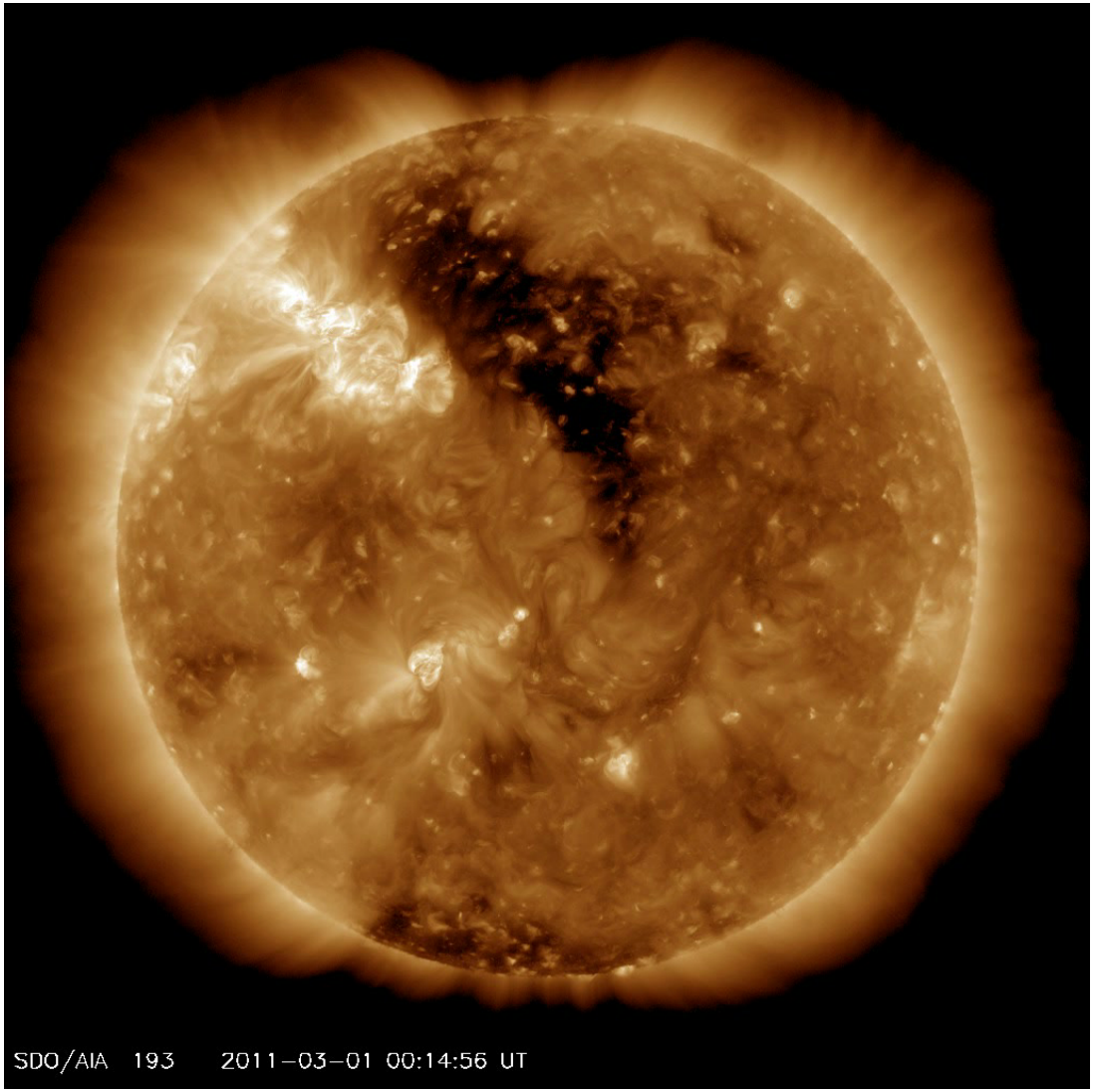


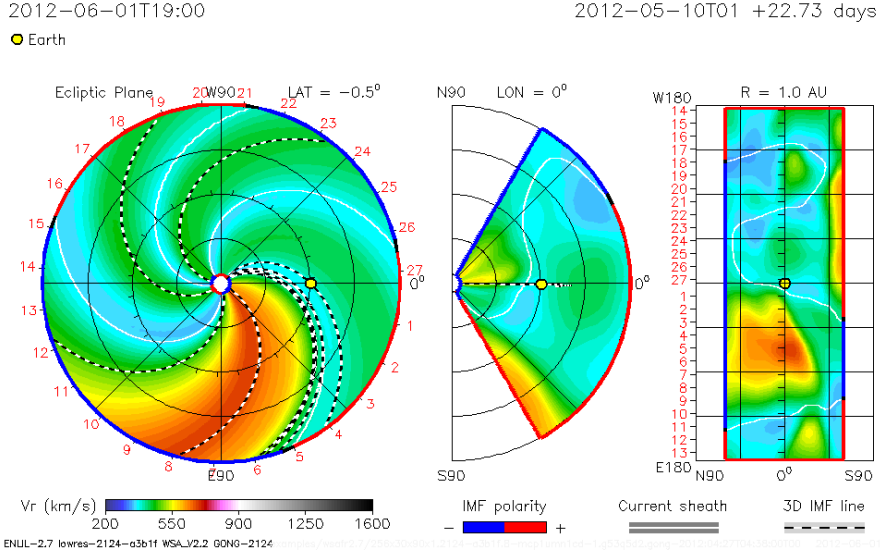
FIG. 4.13. Geometry of the interaction between fast solar wind (on less tightly wound spiral streamlines) and ambient solar wind (on more tightly wound spiral streamlines). The plasma is compressed where streamlines converge. (From Pizzo, 1985.)

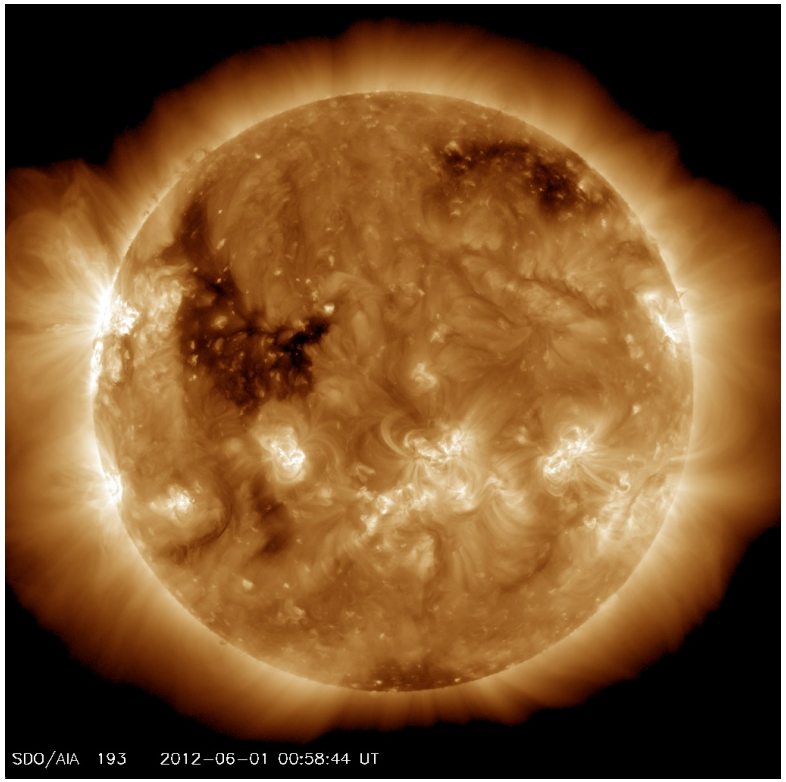
Coronal Hole HSS



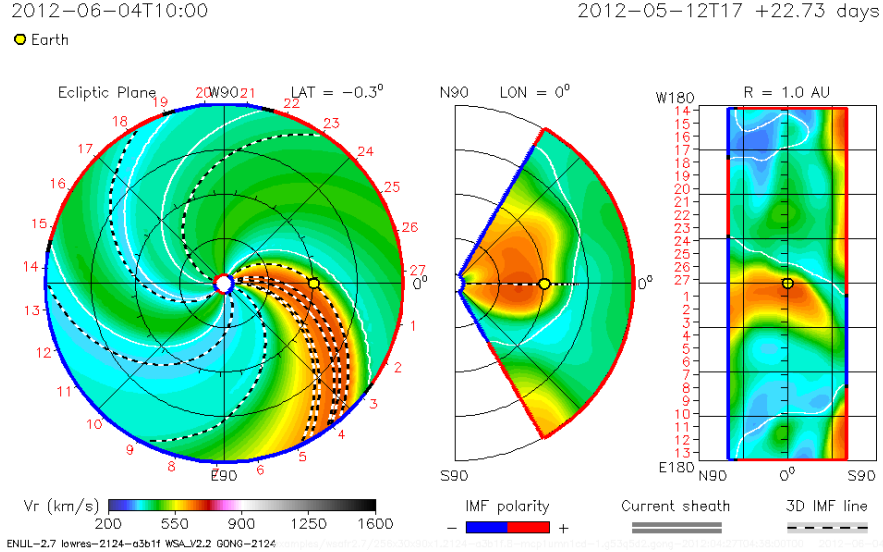
Mar 1, 2011

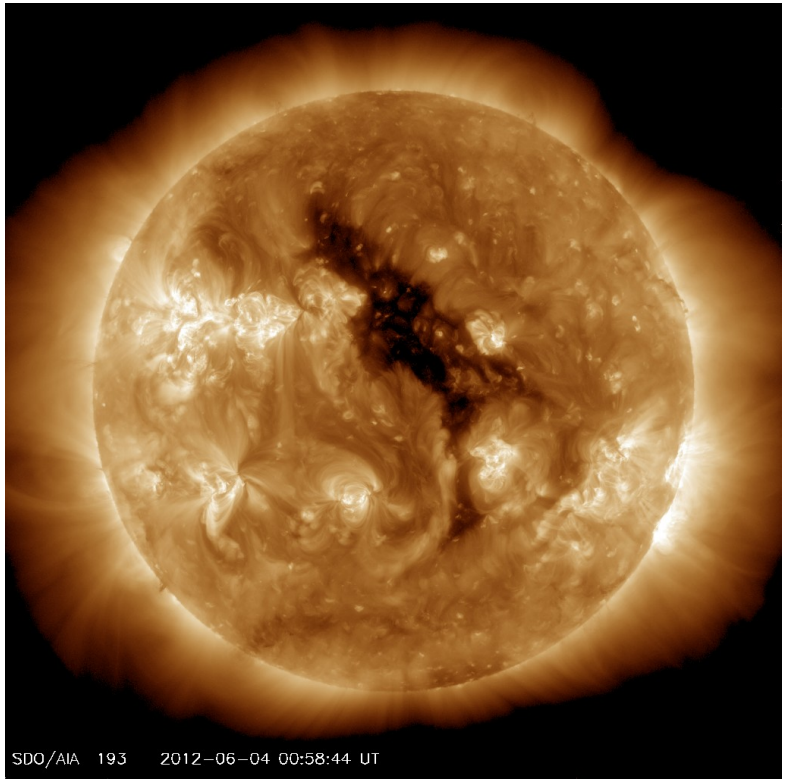
2012-06-01T19:00 2012-05-10T01 +22.73 days





2012-06-04T10:00 2012-05-12T17 +22.73 days



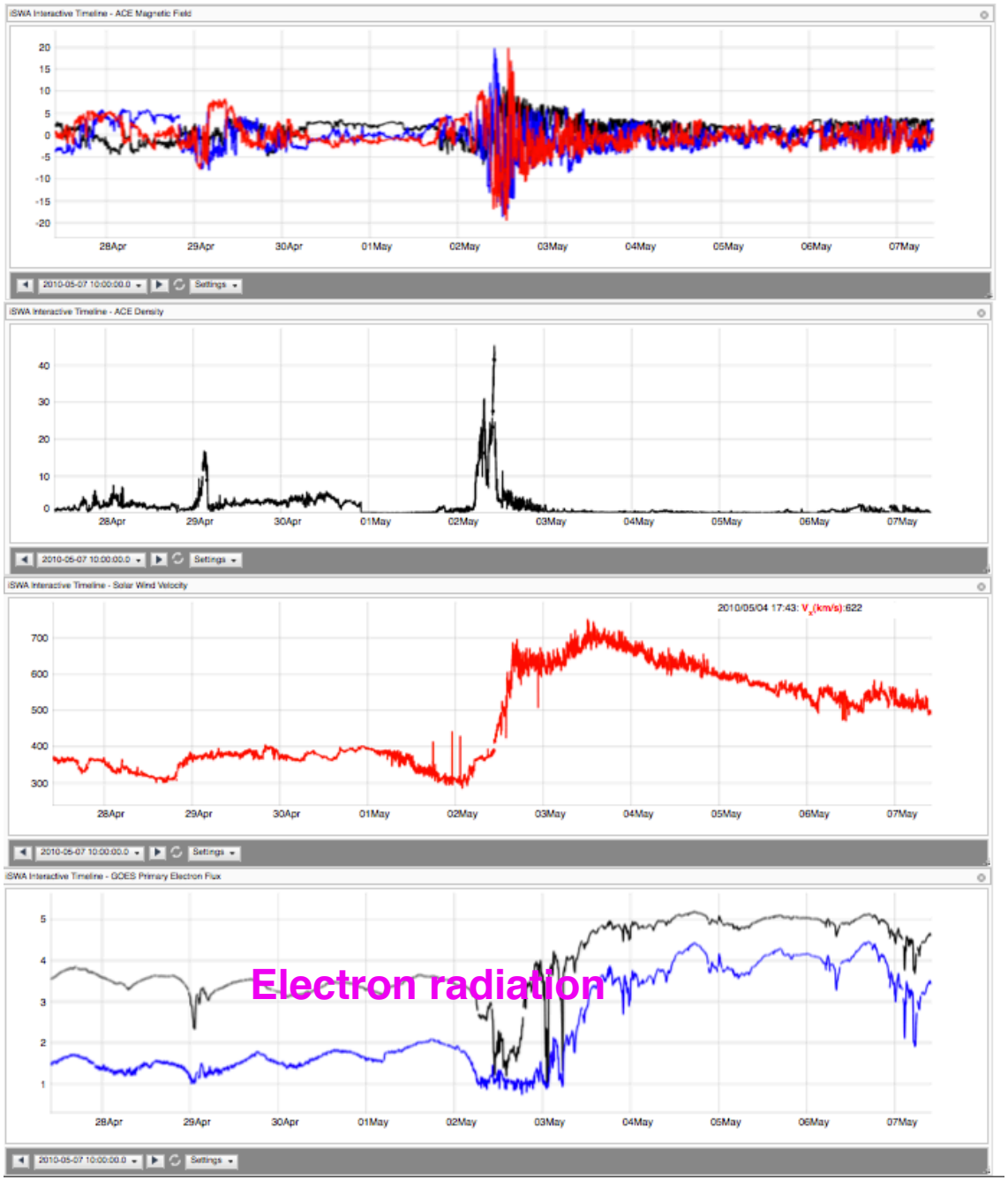


Coronal Hole HSS

- ✓ Is one IMPORTANT space weather contributor too!
- ✓ Particularly for its role in enhancing electron radiation levels near **GEO** orbit and for substantial energy input into the Earth's upper atmosphere
- ✓ May be more hazardous to Earth-orbiting satellites than CME-related magnetic storm particles and solar energetic particles (SEP)

In-situ signatures of CME and CIR HSS at L1

Observing satellites: ACE and WIND



Clean HSS

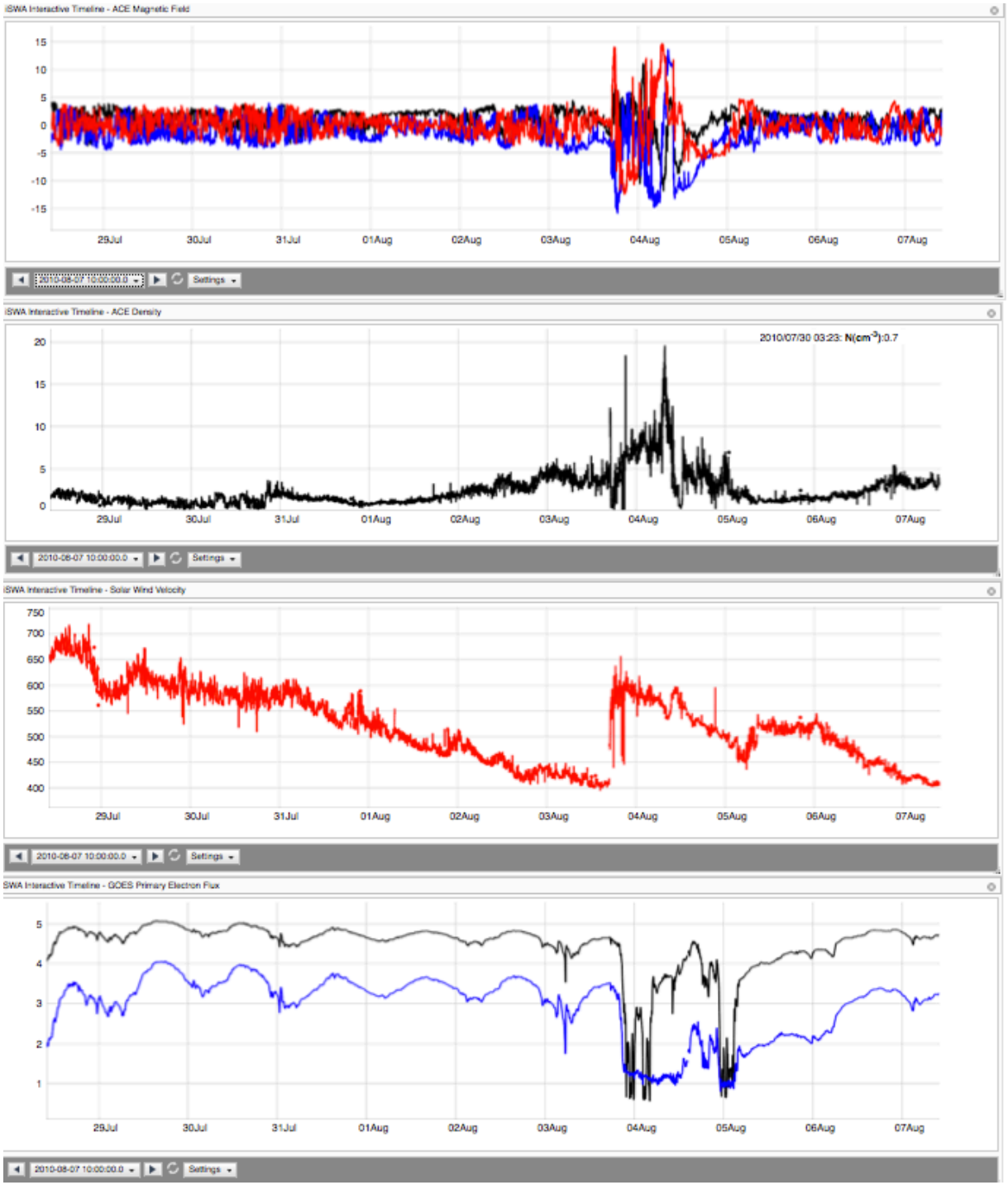
May 2, 2010

Dense (20-30 cc), HSS

IMFBz: -18 nT

may be more hazardous to Earth-orbiting satellites than ICME-related magnetic storm particles and solar energetic particles

Aug 3, 2010

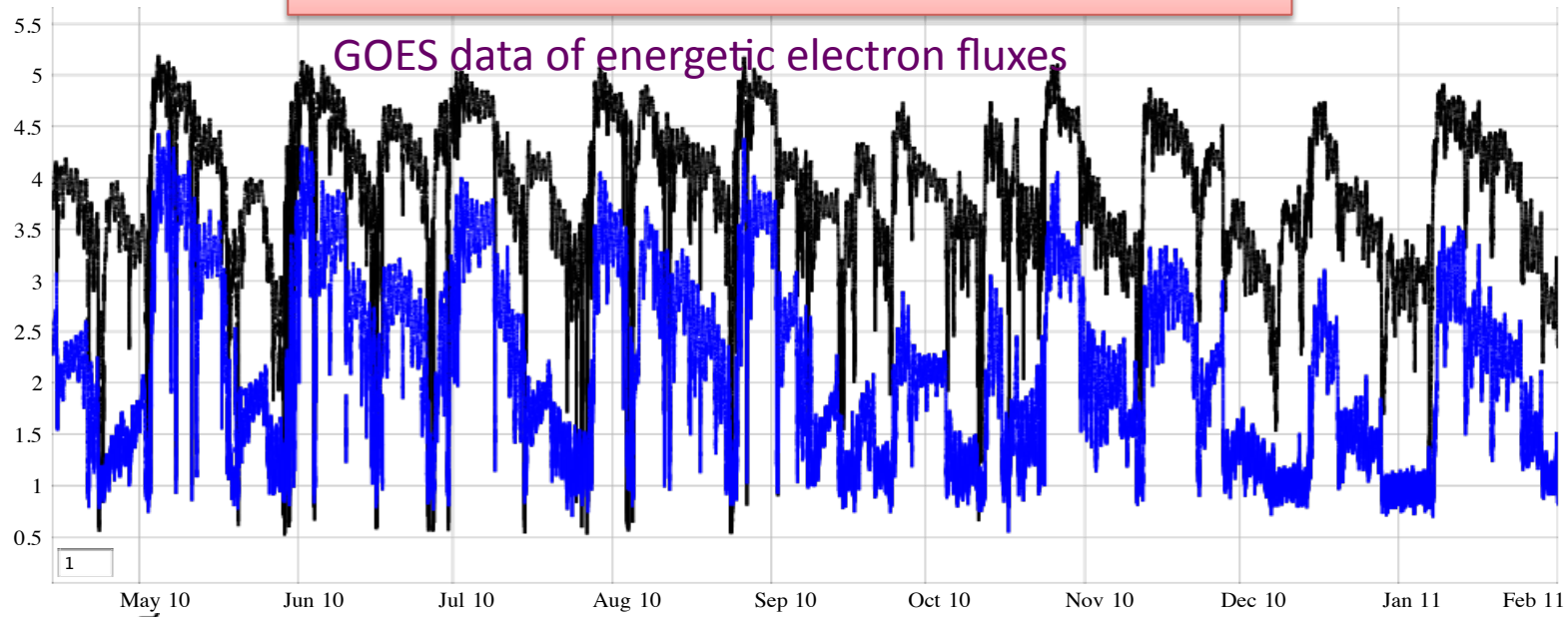

ISWA Interactive Timeline - ACE Magnetic Field
15
10
5
0
-5
-10
-15
29Jul
30Jul
31Jul
01Aug
02Aug
03Aug
04Aug
05Aug
06Aug
07Aug
2010-08-07 10:00:00.0
Settings
ISWA Interactive Timeline - ACE Density
2010/07/30 03:23: N(cm-3):0.7
20
15
10
5
0
29Jul
30Jul
31Jul
01Aug
02Aug
03Aug
04Aug
05Aug
06Aug
07Aug
2010-08-07 10:00:00.0
Settings
ISWA Interactive Timeline - Solar Wind Velocity
750
700
650
600
550
500
450
400
29Jul
30Jul
31Jul
01Aug
02Aug
03Aug
04Aug
05Aug
06Aug
07Aug
2010-08-07 10:00:00.0
Settings
SWA Interactive Timeline - GOES Primary Electron Flux
5
4
3
2
1
29Jul
30Jul
31Jul
01Aug
02Aug
03Aug
04Aug
05Aug
06Aug
07Aug
2010-08-07 10:00:00.0
Settings

Different impacts on RB
CME vs CIR storms

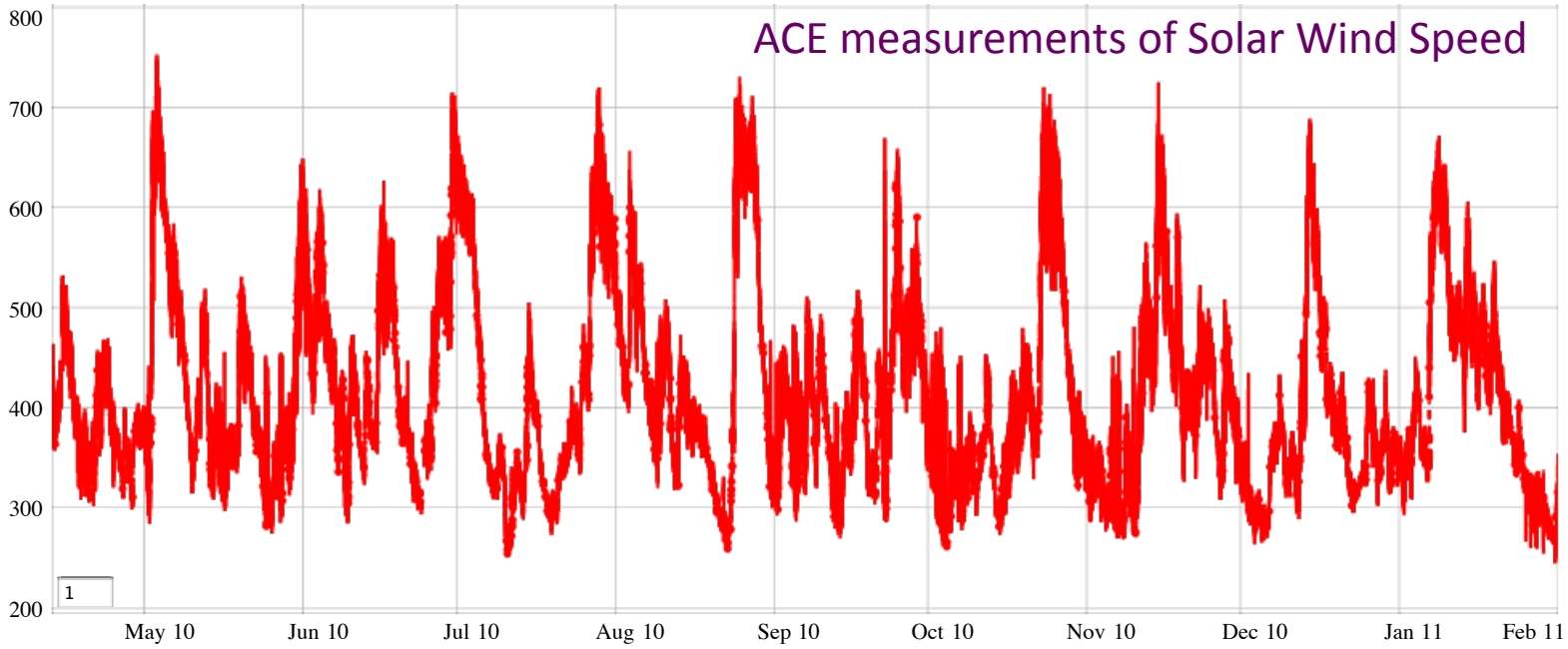
- CME geomagnetic storms: RB flux peak inside geosynchronous orbit. The peak locations moves inward as storm intensity increases
- CIR geomagnetic storms: More responsible for the electron radiation level enhancement at GEO orbit

Click the check boxes to

HSS and radiation belt electron flux enhancement

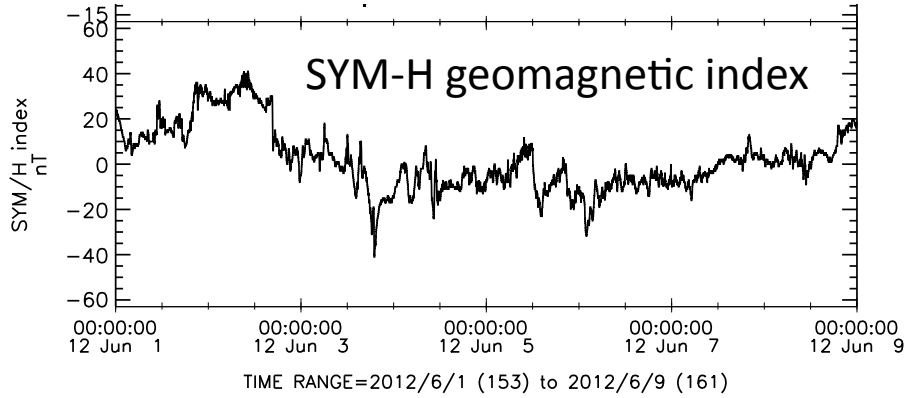


☑E > 0.8 MeV ☑E > 2.0 MeV **Zoom:** In Out full **Pan:** left right



☑Bulk Speed **Zoom:** In Out full **Pan:** left right

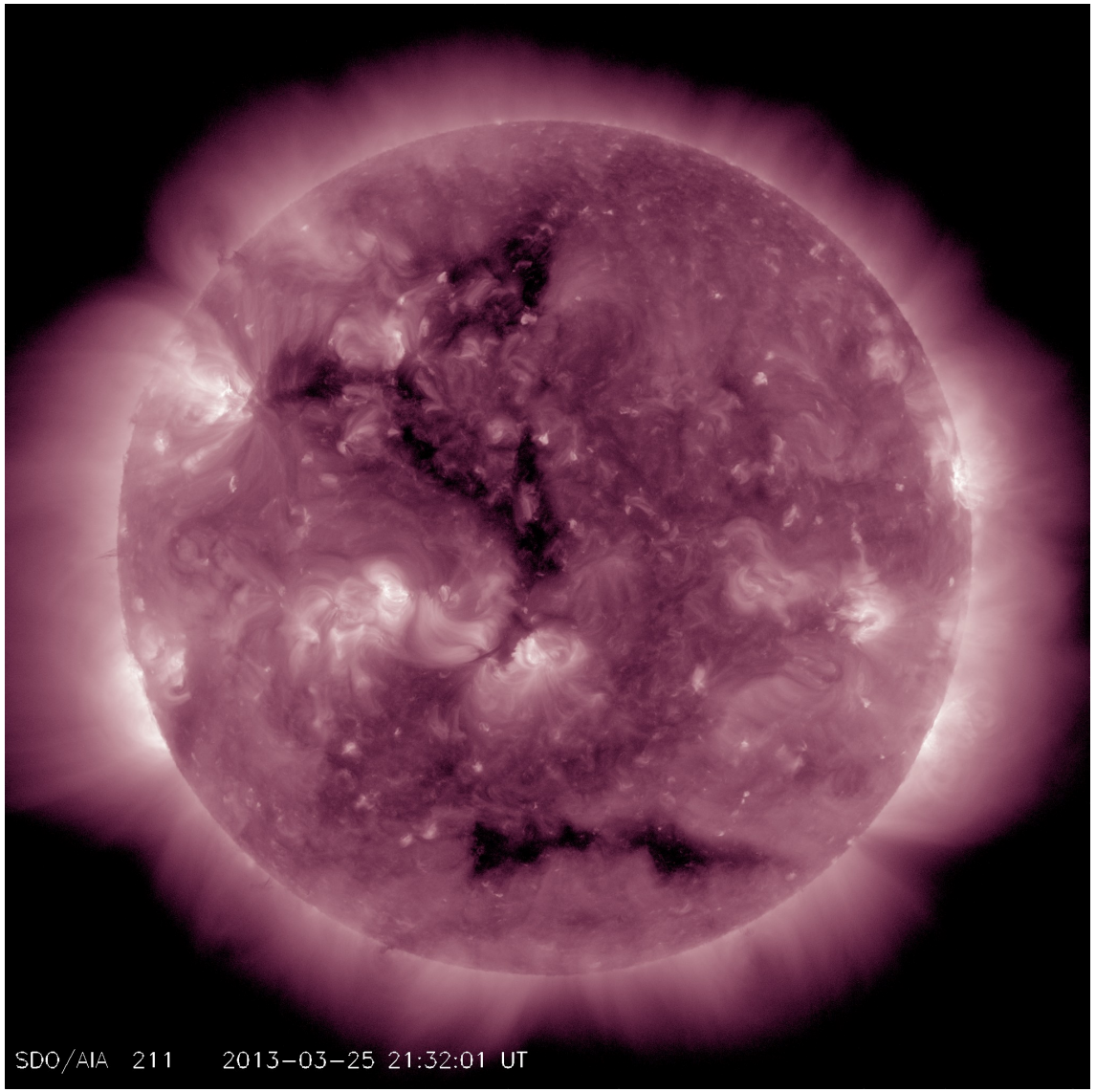
- The increase in speed from a solar wind high speed stream pumps energy into the magnetosphere which can cause **geomagnetic storms** and **energizes particles**.
- They can produce energetic **electron flux enhancements** in the radiation belt.
- Geomagnetic storms are disturbances/changes in Earth's magnetic field due to changes in solar wind conditions typically lasting 3-6 days.
- High speed streams can also cause geomagnetic storms, however they are longer in duration and not as strong as geomagnetic storms caused by CMEs.



Coronal Hole High Speed Streams can have ‘far-reaching’ effects

10 AU at Saturn

The Coronal Hole possibly responsible for the aurora at Saturn around May

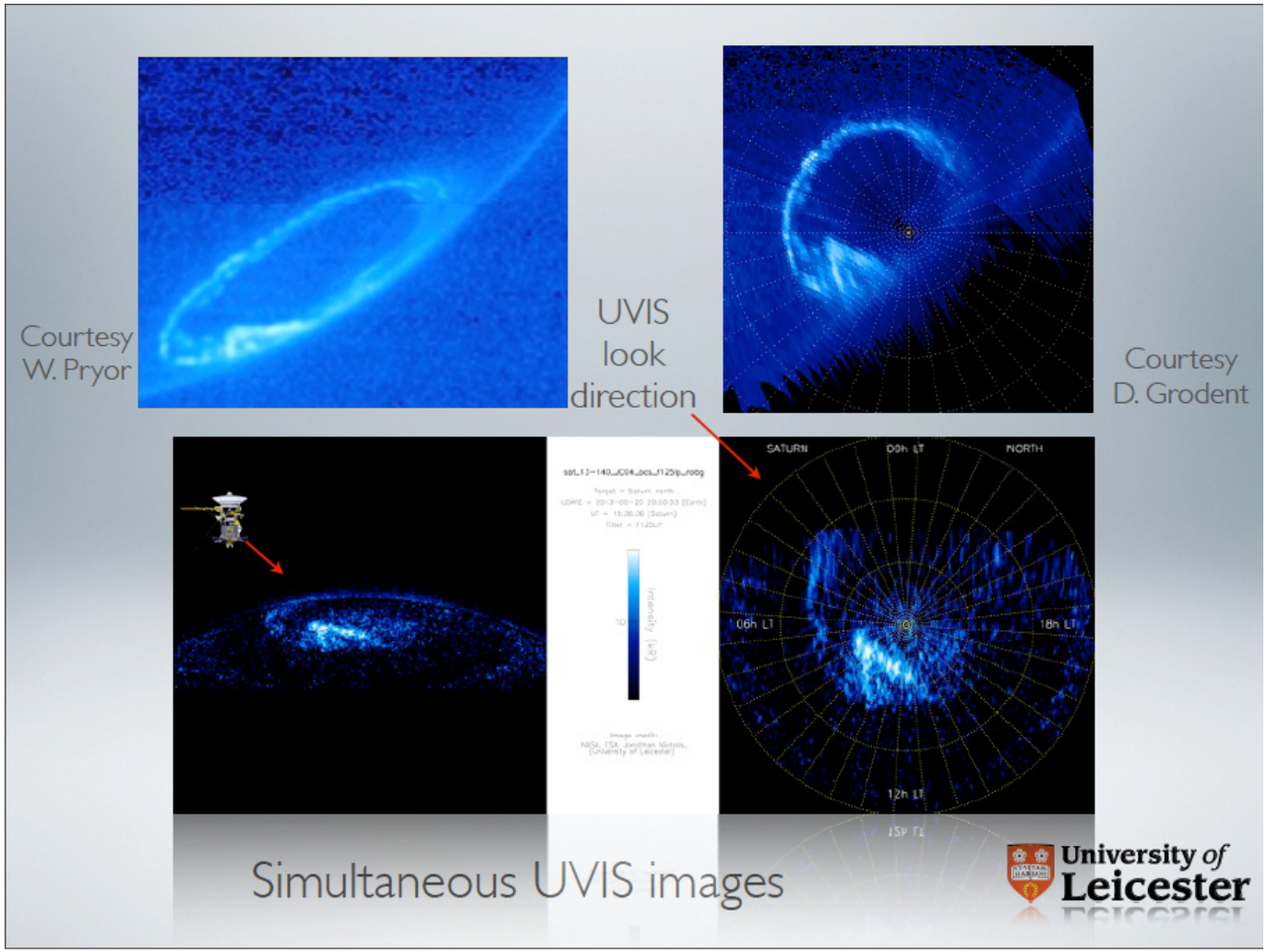


Near simultaneous CASSNI/UVIS and HST images

(J. Nichols)

Raw UVIS

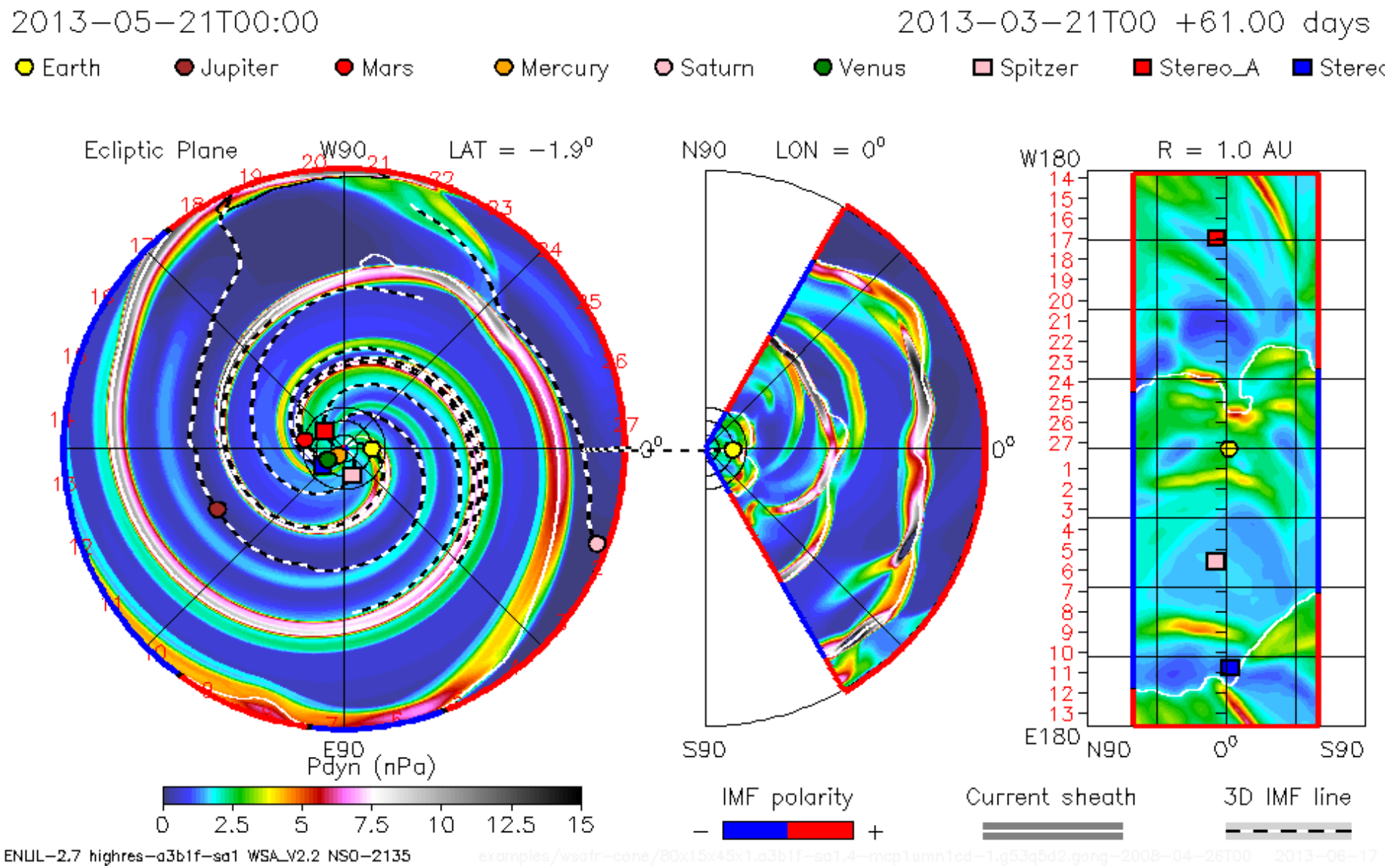
Raw HST



Polar UVIS

Polar HST

WSA+ENLIL Simulation Result for Saturn



Coronal Holes and High Speed Streams – Assessment

Fill out the form:
http://bit.ly/swassess2

Why is there solar wind?

Does slow solar wind generally originate from high or low latitudes on the Sun?
What is the typical speed for slow solar wind?
What is the typical density for slow solar wind?

Does fast solar wind generally originate from high or low latitudes on the Sun?
What is the typical speed for fast solar wind?
What is the typical density for fast solar wind?

What is a coronal hole? How does the plasma density of a coronal hole compare to the rest of the corona?

Where are coronal holes typically found during solar minimum, and how does that change at solar maximum?

What are some wavelengths we can view the sun

What is a high speed stream, and how is it related to a coronal hole?

What is a corotating interaction region (CIR) and stream interaction region (SIR)?

How does a high speed stream caused geomagnetic storm differ from one caused by a CME?

How are high speed streams related to energetic electron fluxes in the magnetosphere?

Update the time on various cygnets on this layout to show one high speed stream impacting Earth:
http://go.nasa.gov/15yyWcP

Did the >0.8 MeV energetic electron flux measured at GOES increase due to your chosen high speed stream? Did it go above the alert threshold value of 10^5 pfu? If so, at what time? http://go.nasa.gov/11pZxbR

Slide link summary

SW REDI website
http://ccmc.gsfc.nasa.gov/support/SWREDI/swredi.php
iSWA http://iswa.gsfc.nasa.gov
iSWA Cygnet Glossary http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Full_iSWA_Cygnet_List
iSWA Space Weather Glossary http://iswa3.ccmc.gsfc.nasa.gov/wiki/index.php/Glossary

Most figures and tables in the slides are from the "Introduction to Space Physics" textbook
http://www.cambridge.org/us/knowledge/isbn/item1145043

iSWA layouts of high speed streams http://go.nasa.gov/17nkicp
http://1.usa.gov/1avidfv (coronal hole and in-situ)
A Coronal hole in field view of SDO and STEREO A, B
http://1.usa.gov/LSfnaC (March 2013)

Coronal hole example video:
http://www.youtube.com/watch?v=xNISXItrPLg&feature=youtube_gdata

Saturn aurora
1. an iSWA layout for the coronal hole (temporal evolution) while it was the Earth-facing solar disk
http://1.usa.gov/1fk5Fv6

2) SDO view of the coronal hole when it arrived at ACE and the ACE in-situ signatures of the CIR/HSS
http://1.usa.gov/1lZjWeD

Extras

WSA+ENLIL+Cone for Saturn